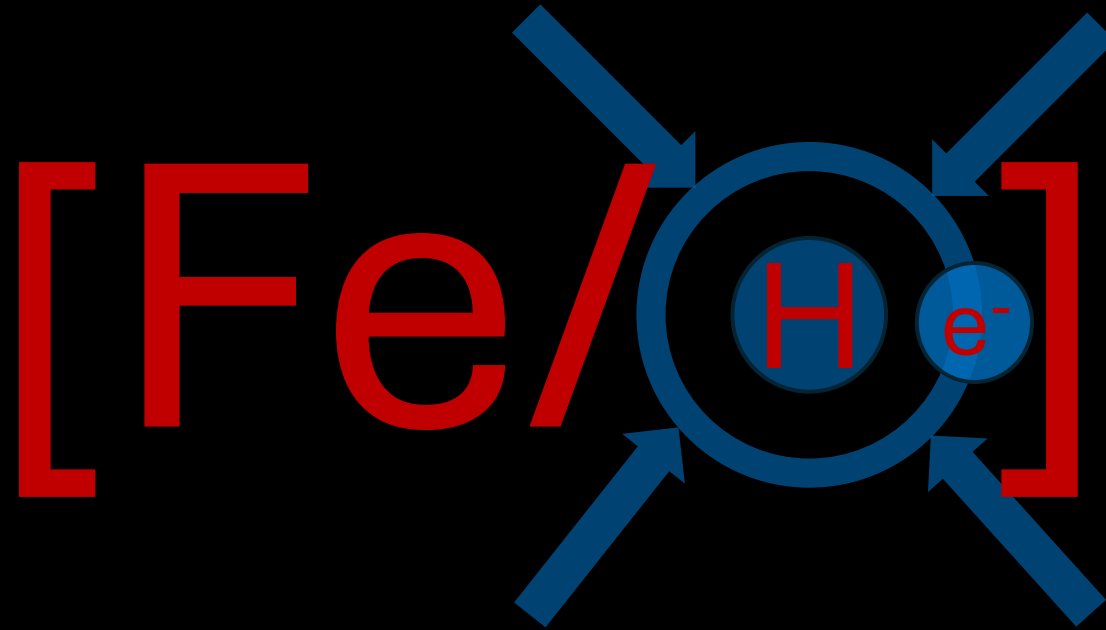


Why Galactic Chemical Evolution is Actually About Hydrogen

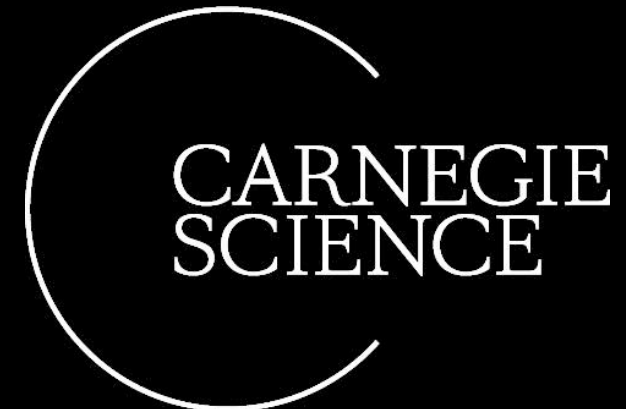


James W. Johnson

Galactic Metal Cycles

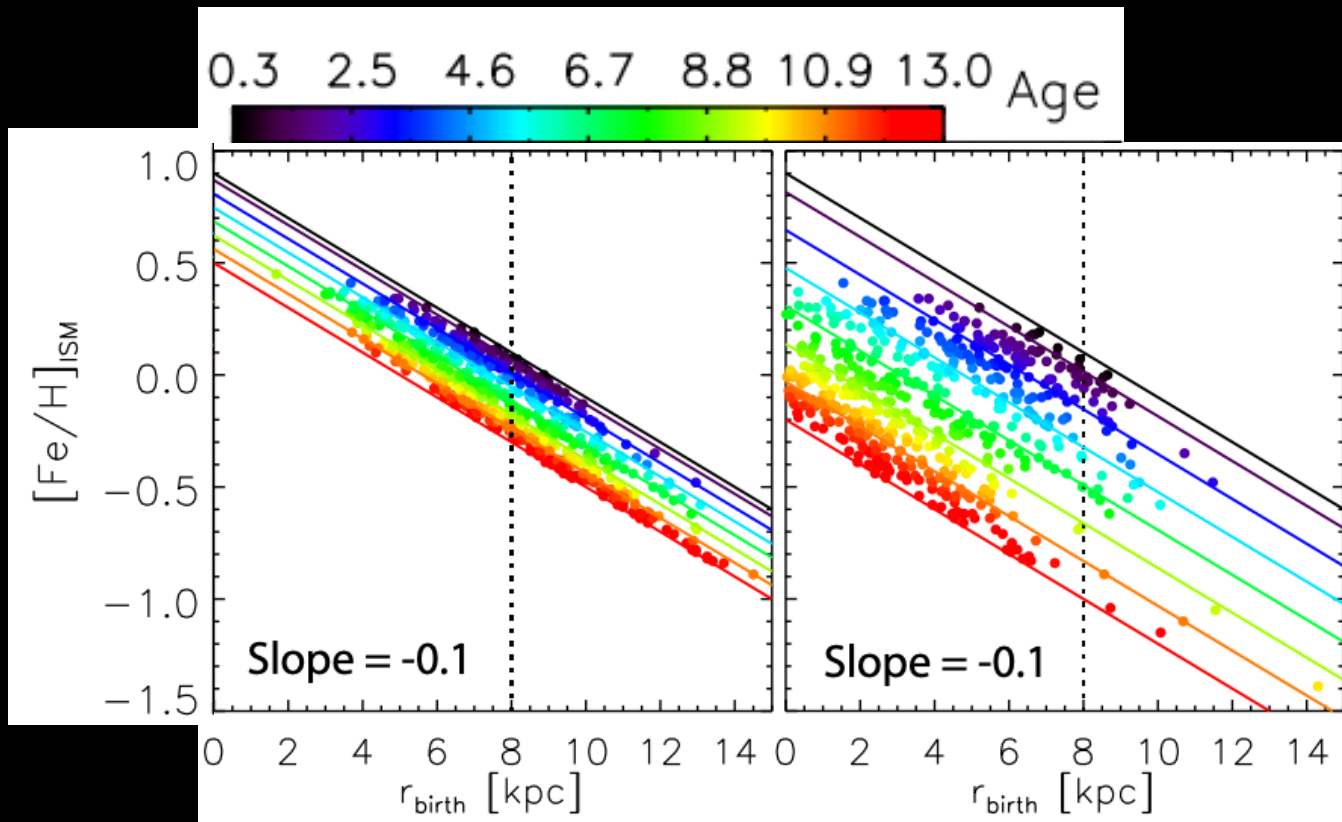
Kylemore Abbey

May 12, 2026



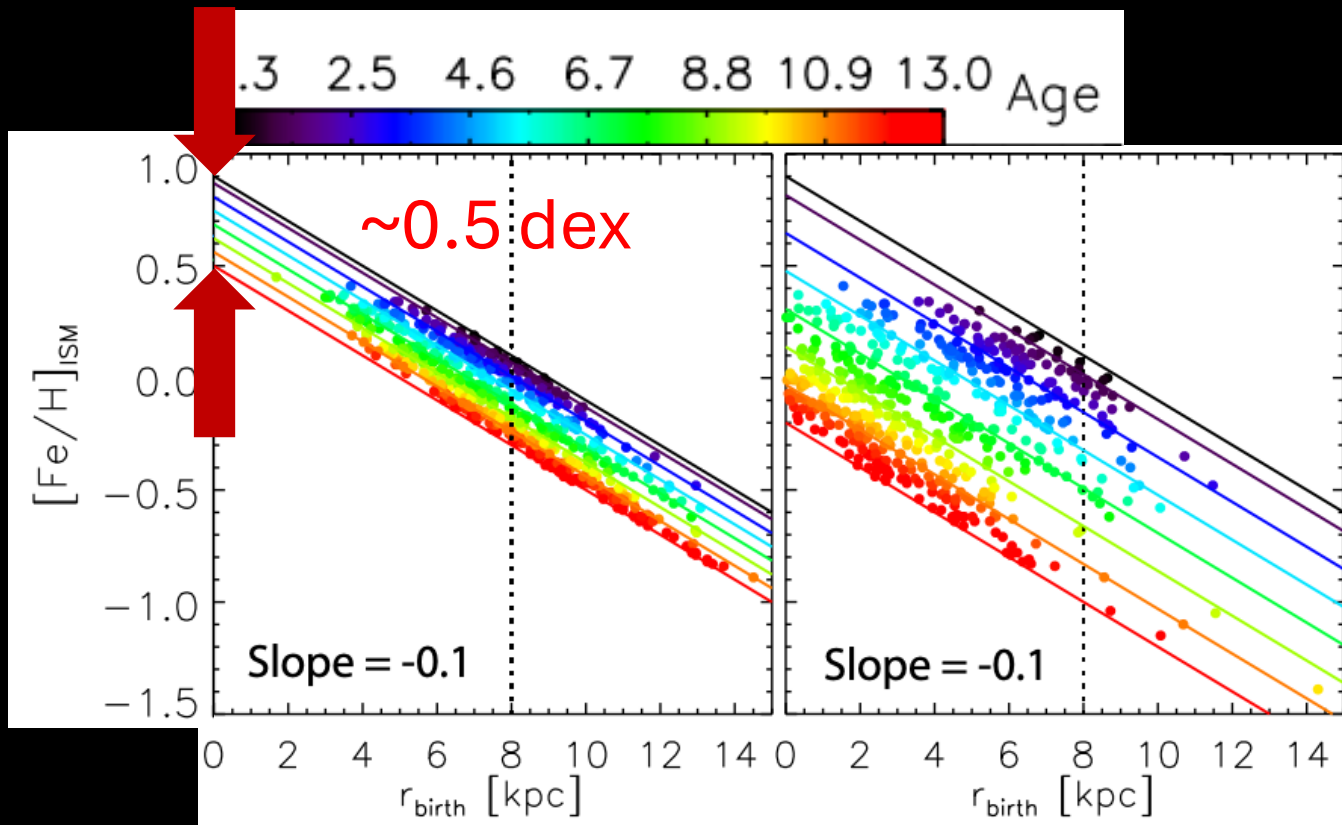
A Conundrum

Models



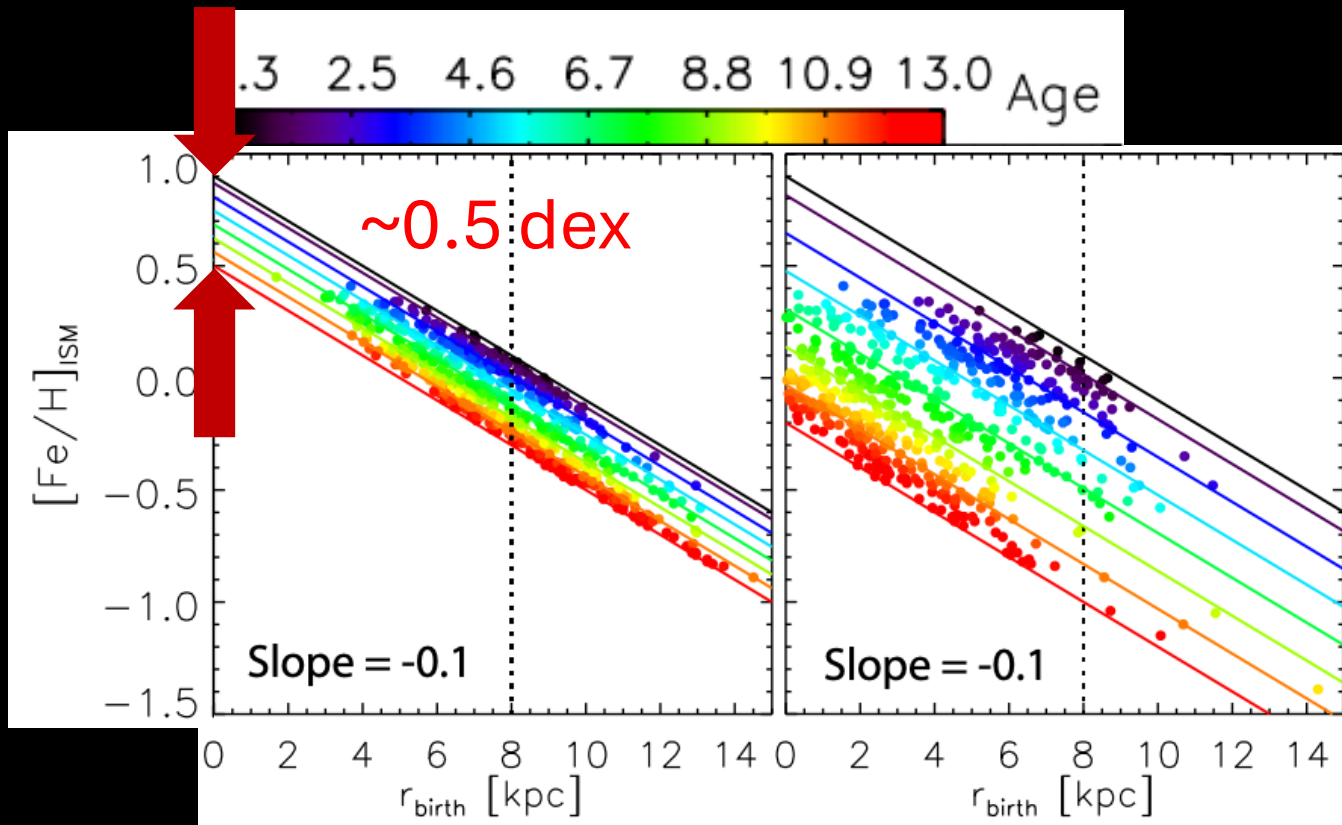
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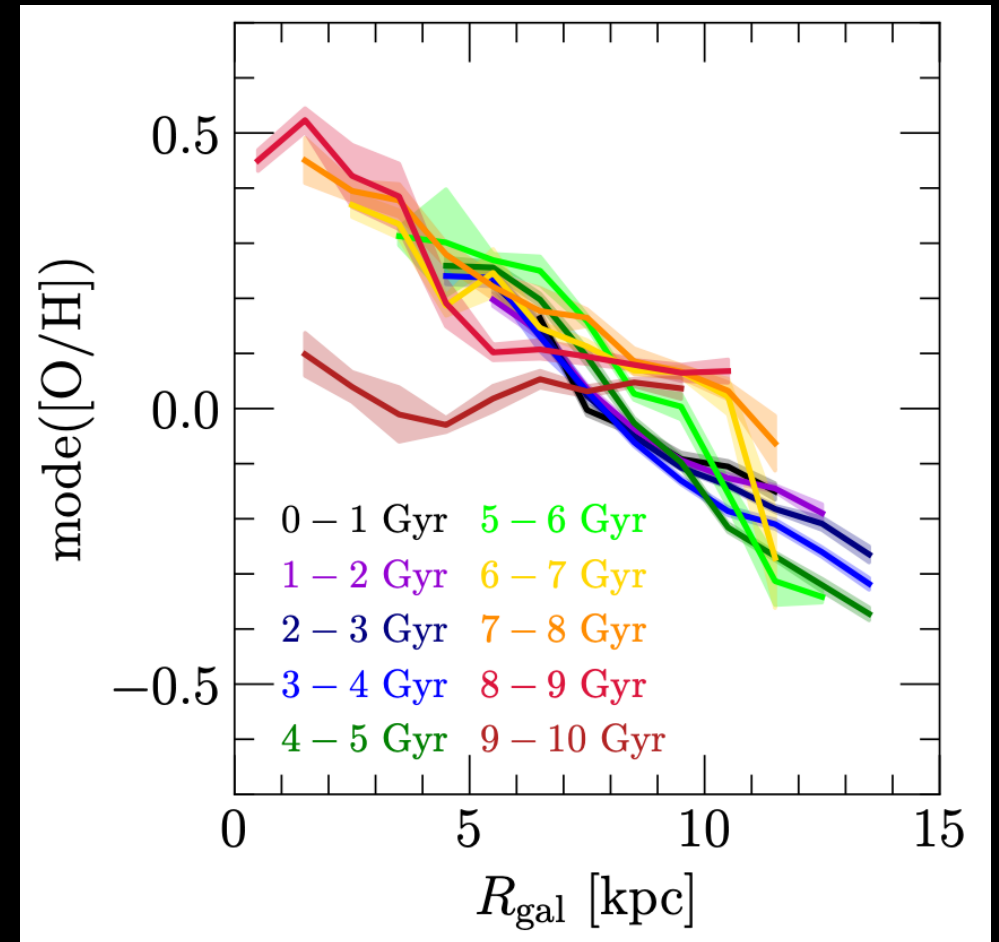


A Conundrum

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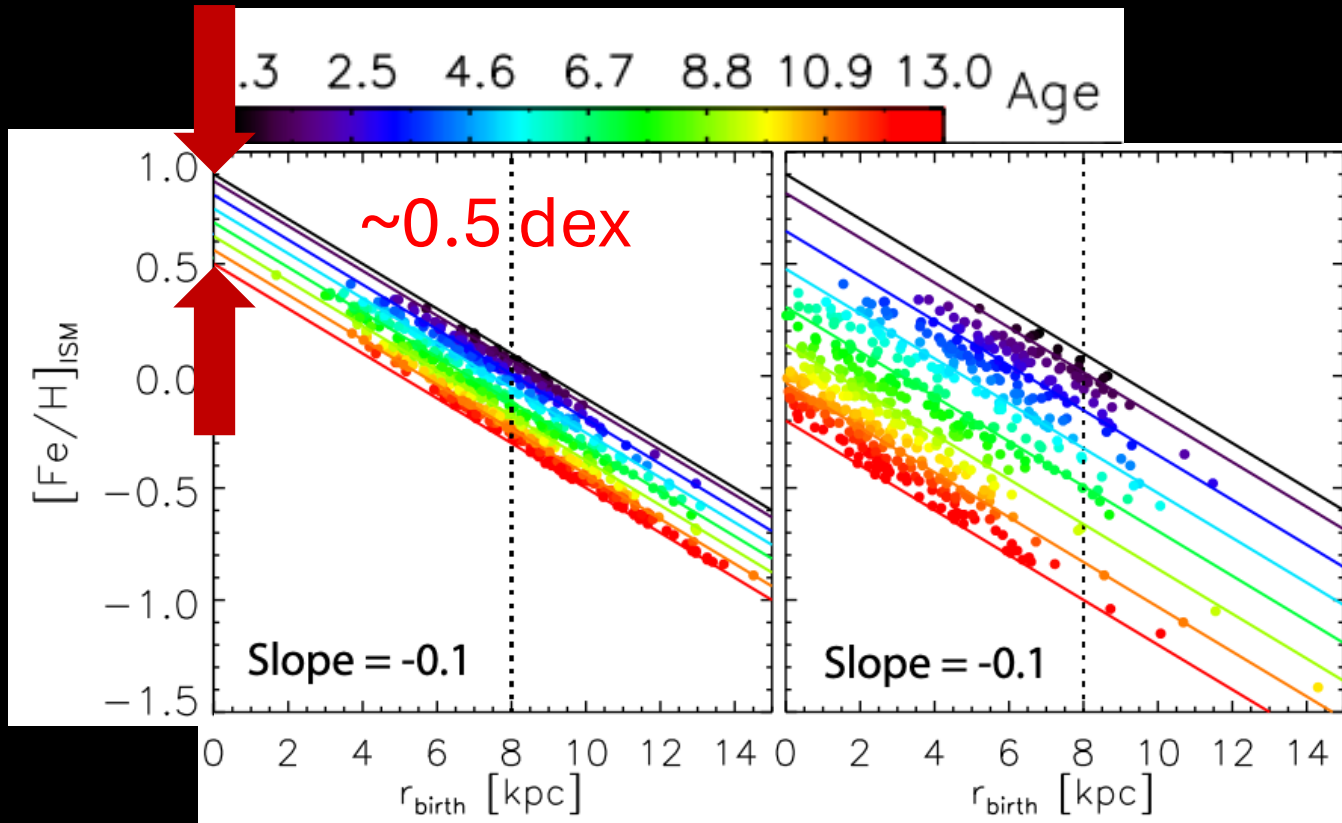


Data

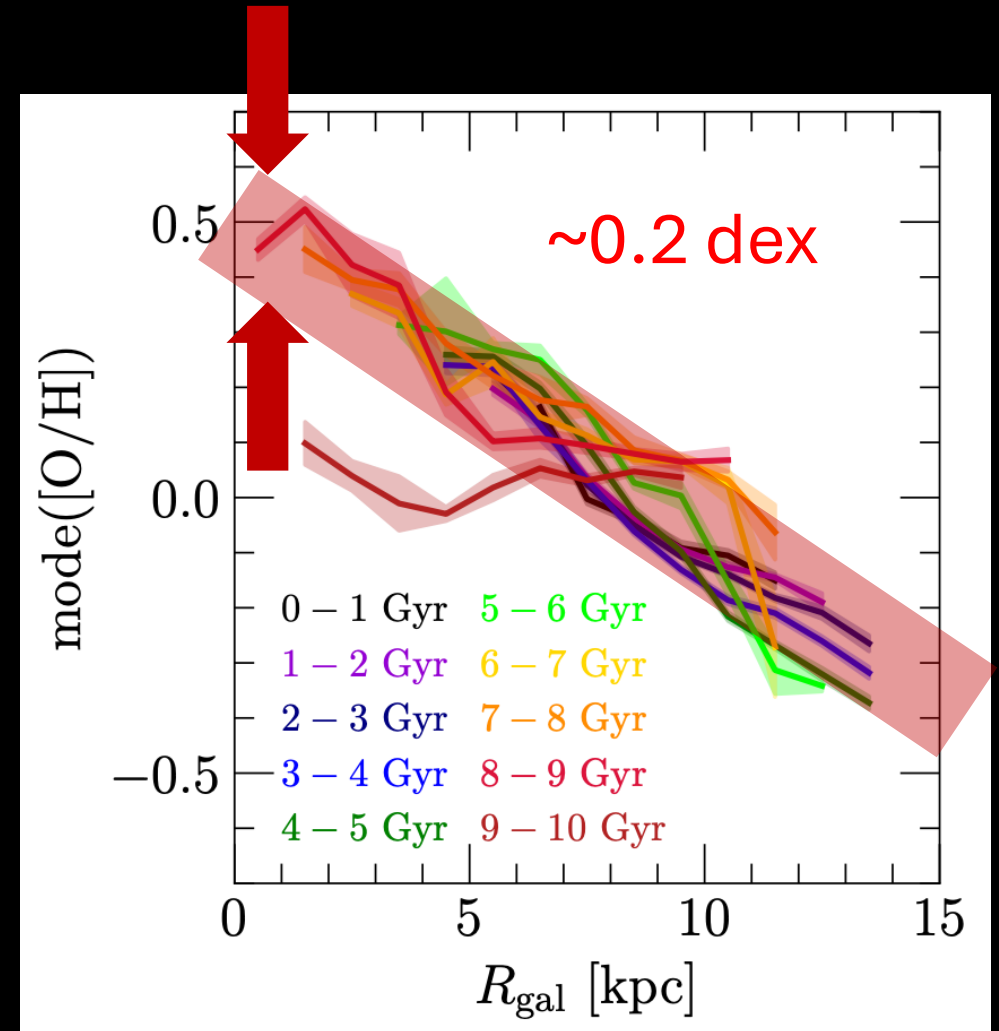


A Conundrum

Models



Data



A History Lesson: “Closed-Box” Models

First GCE models:

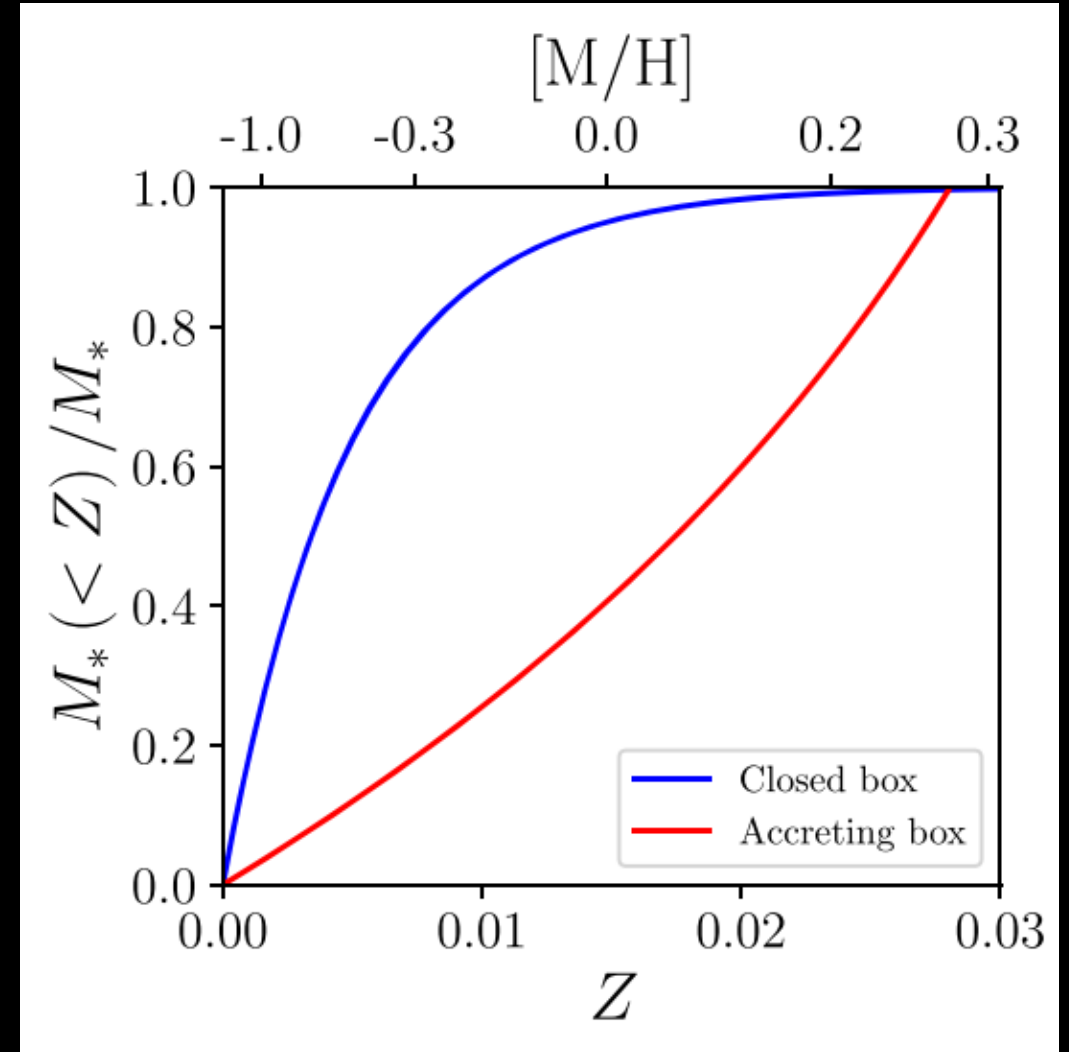
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- Some gas at $t = 0$
- Initially metal-free

A History Lesson: “Closed-Box” Models

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$$P(Z) \propto e^{-Z}$$



A History Lesson: “Closed-Box” Models

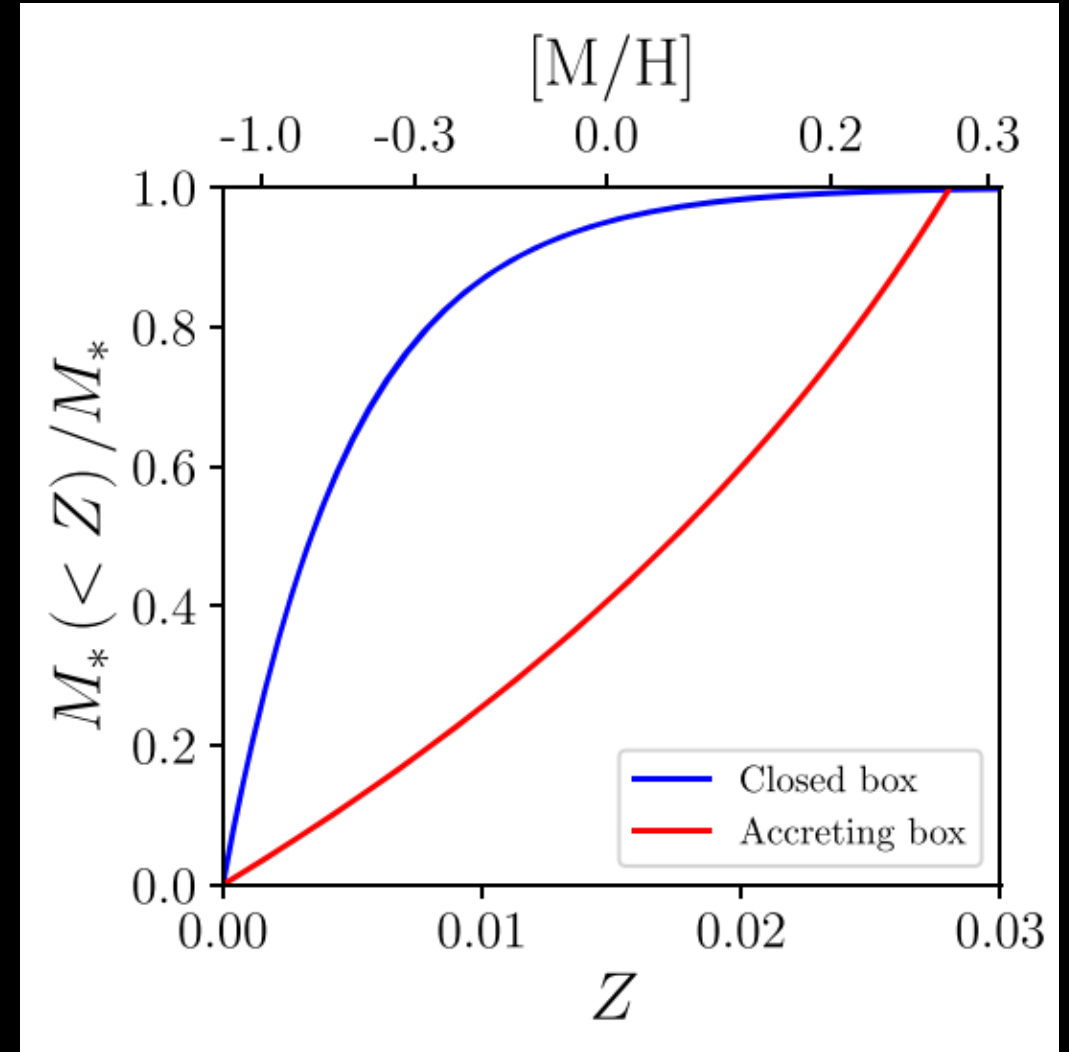
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Solar neighborhood: $P(\ln Z) \sim e^{-(\ln Z)^2}$

- Frequency of low- Z stars overpredicted
- “G-dwarf problem” (Tinsley 1980 review)



A History Lesson: “Closed-Box” Models

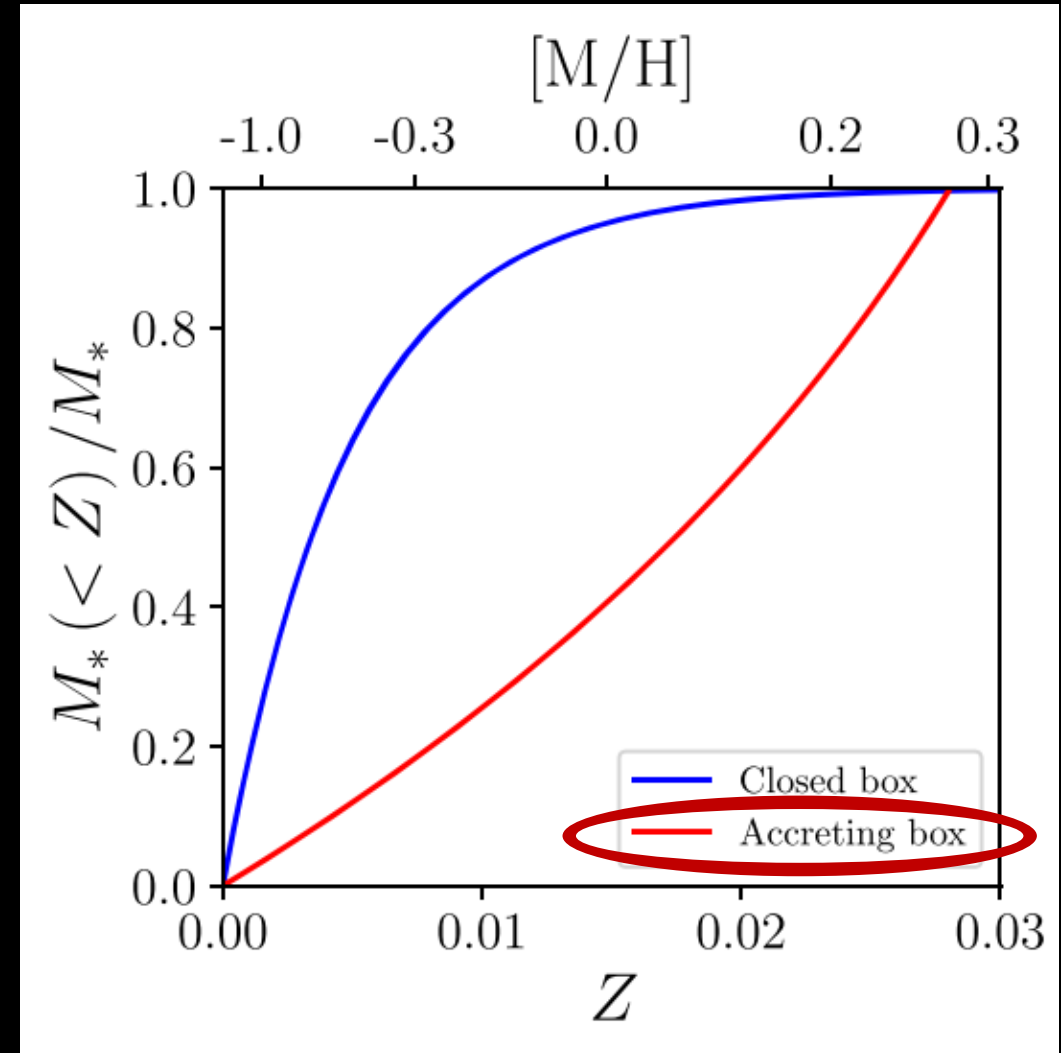
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A History Lesson: The Effect of Accretion

Ensures an ongoing supply of fresh H

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Larson (1972):

$$Z \rightarrow Z_{eq} = \frac{y}{1 - r}$$

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“yield”: mass of metals released
per mass of stars formed

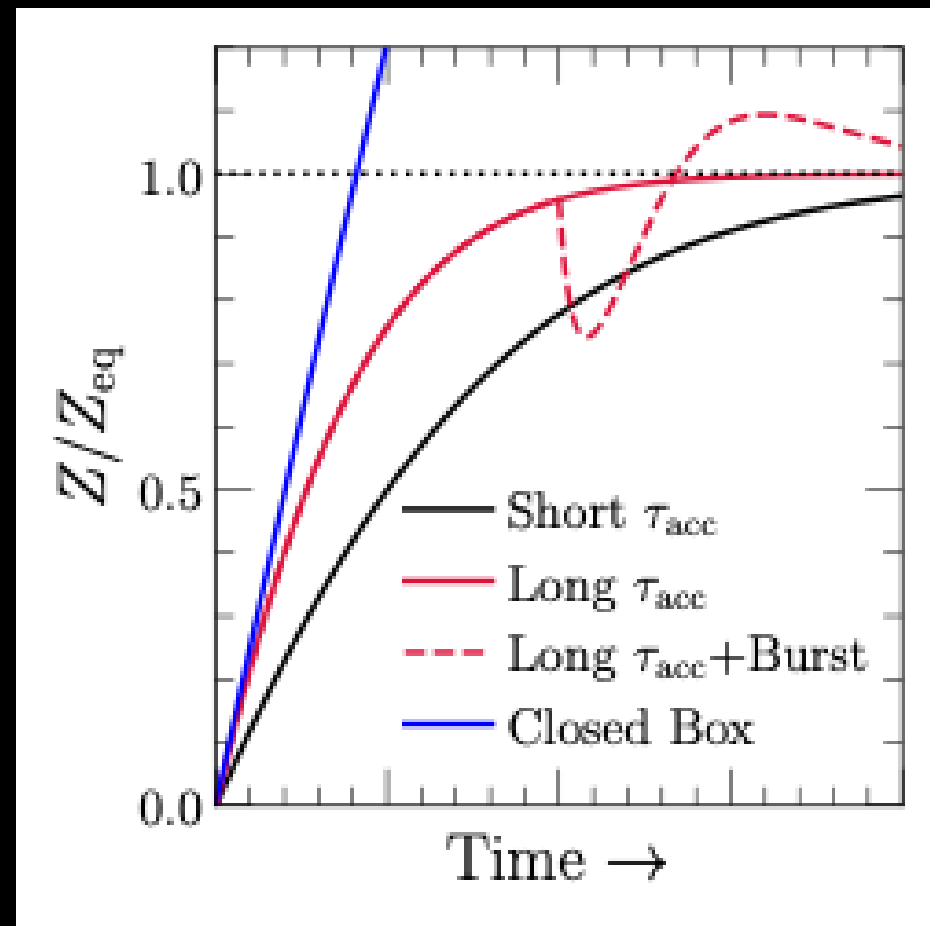
“recycling”: stellar envelopes return
to ISM; ~40% of mass across IMF

A History Lesson: The Effect of Accretion

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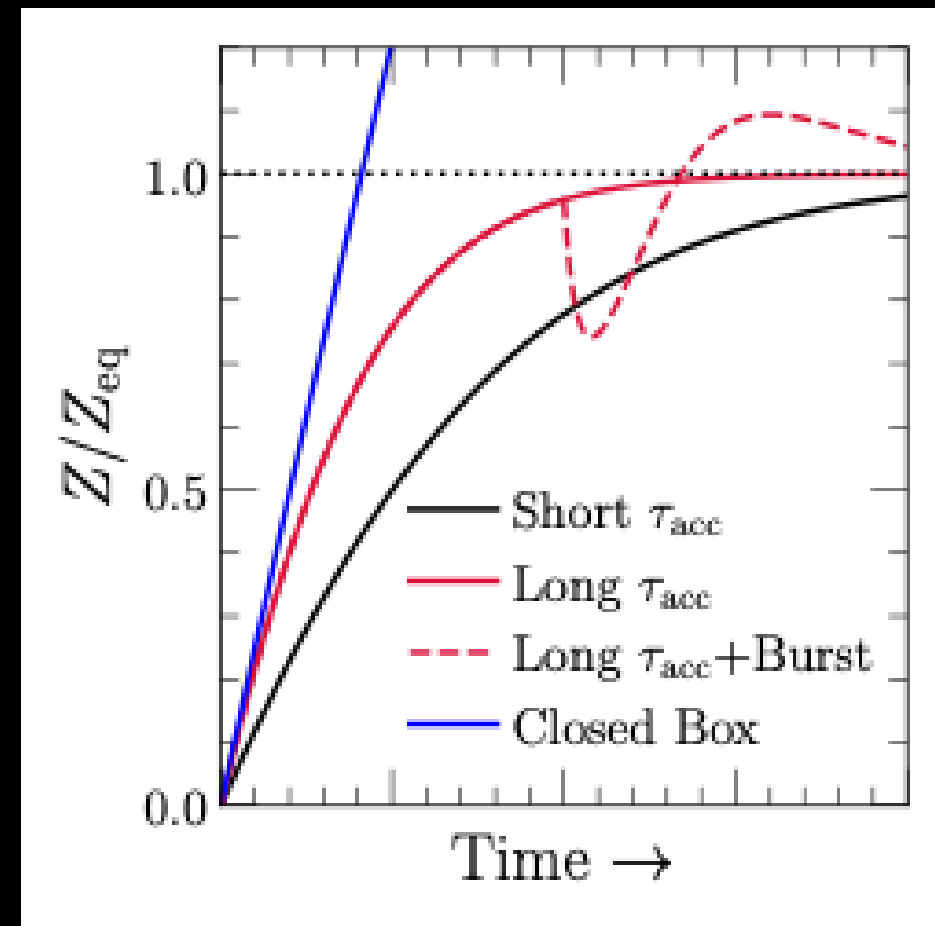
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Closed box: $Z_{eq} \rightarrow \infty$

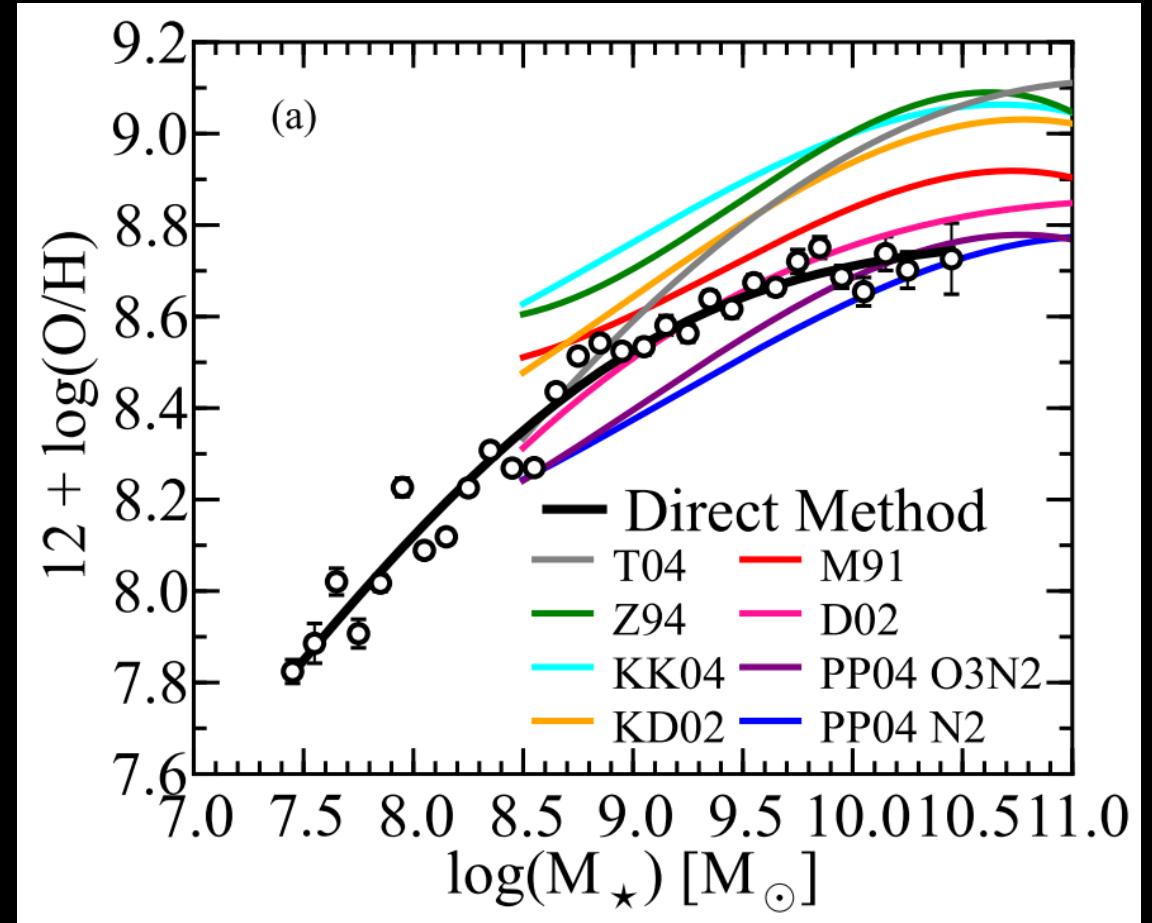
- Eventually all H fuses
- Observationally: $Z_{ISM} \sim Z_{\star}$



A History Lesson: The Effect of Outflows

Mass-metallicity relation for galaxies

- Finlator & Davé (2008)
- Peeples & Shankar (2011)
- Sanchez-Menguiano et al. (2024)



A History Lesson: The Effect of Outflows

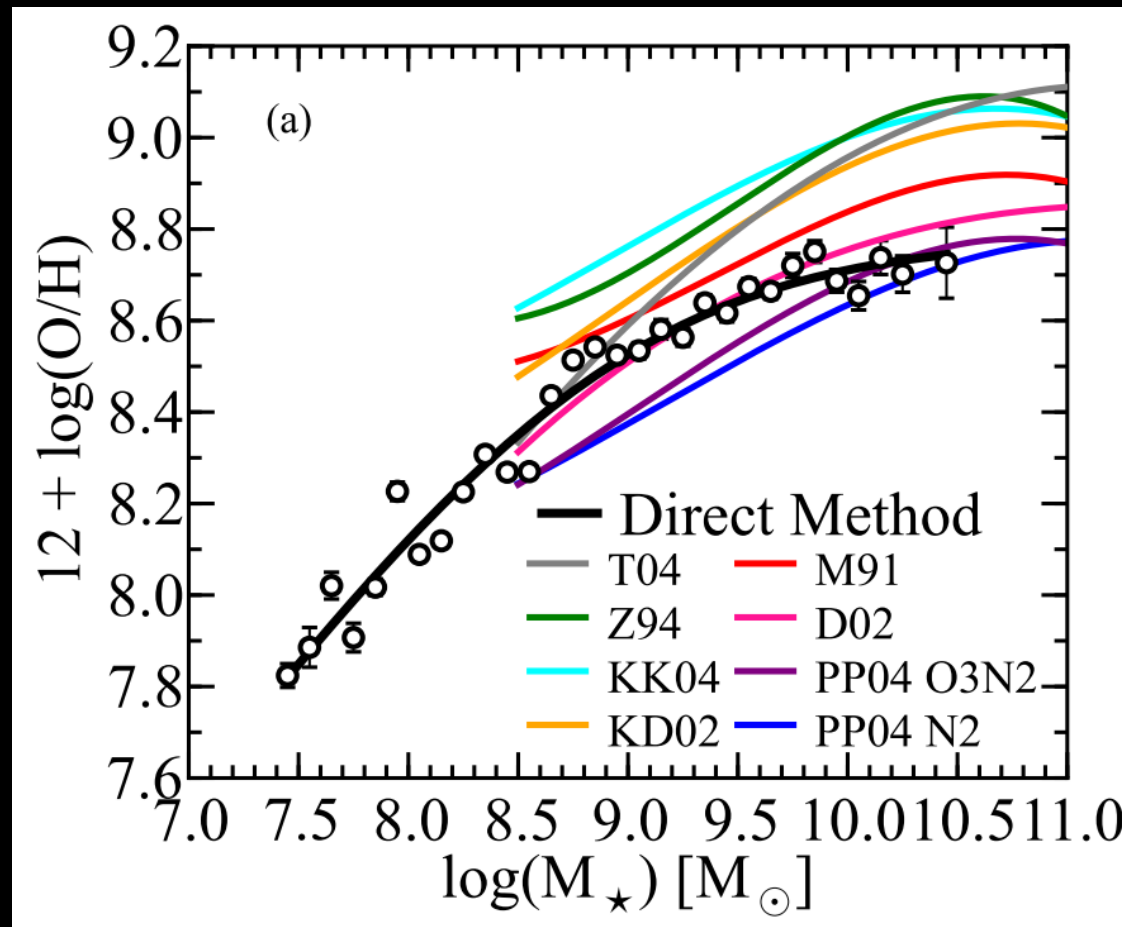
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Lower Z_{eq}

- Need more H to form stars

$$Z \rightarrow Z_{eq} = \frac{y}{1 + \eta - r}$$



A History Lesson: The Effect of Outflows

Mass-metallicity relation for galaxies

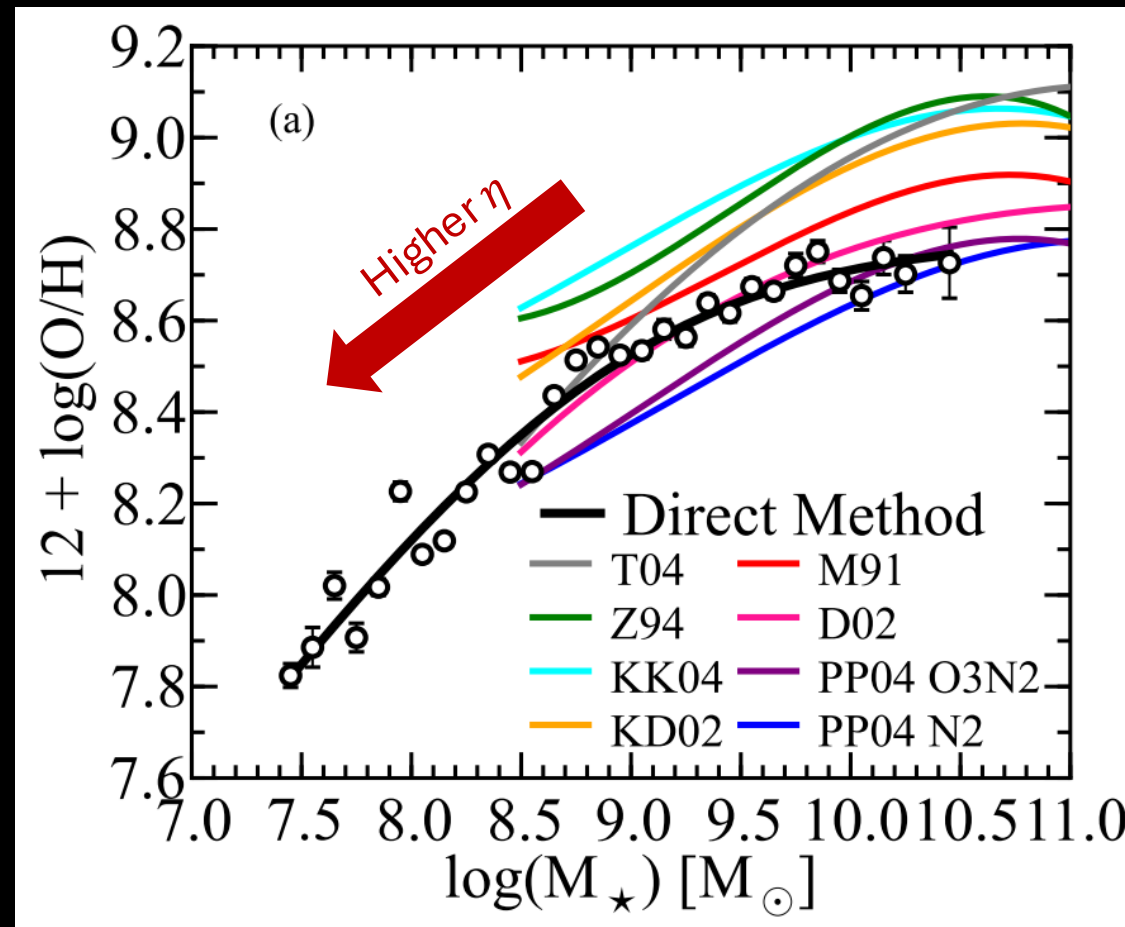
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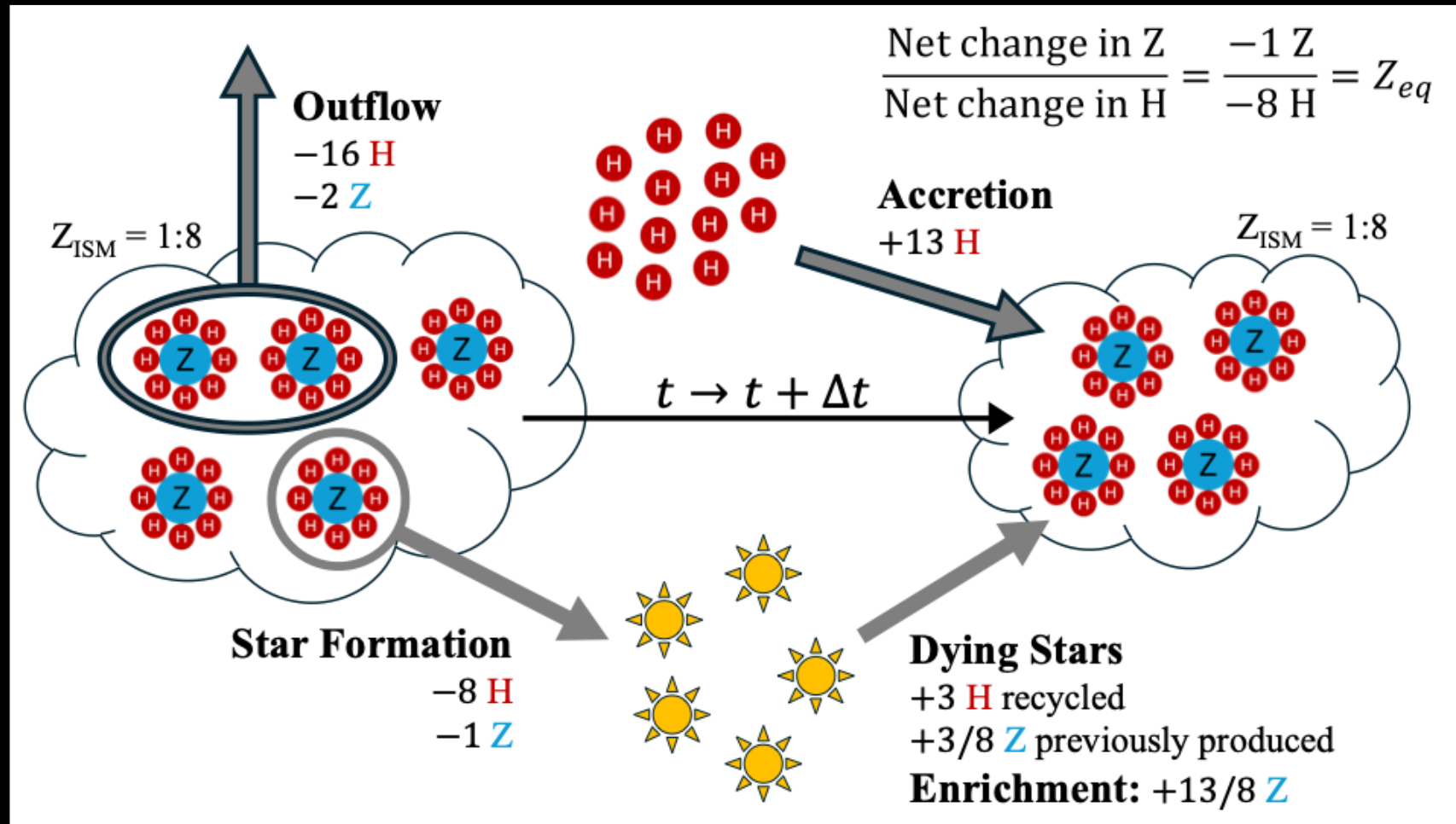
$$Z \rightarrow Z_{eq} = \frac{y}{1 + \eta - r}$$

$$\eta \equiv \dot{M}_{out} / \dot{M}_{\star}$$



Visualizing Chemical Equilibrium

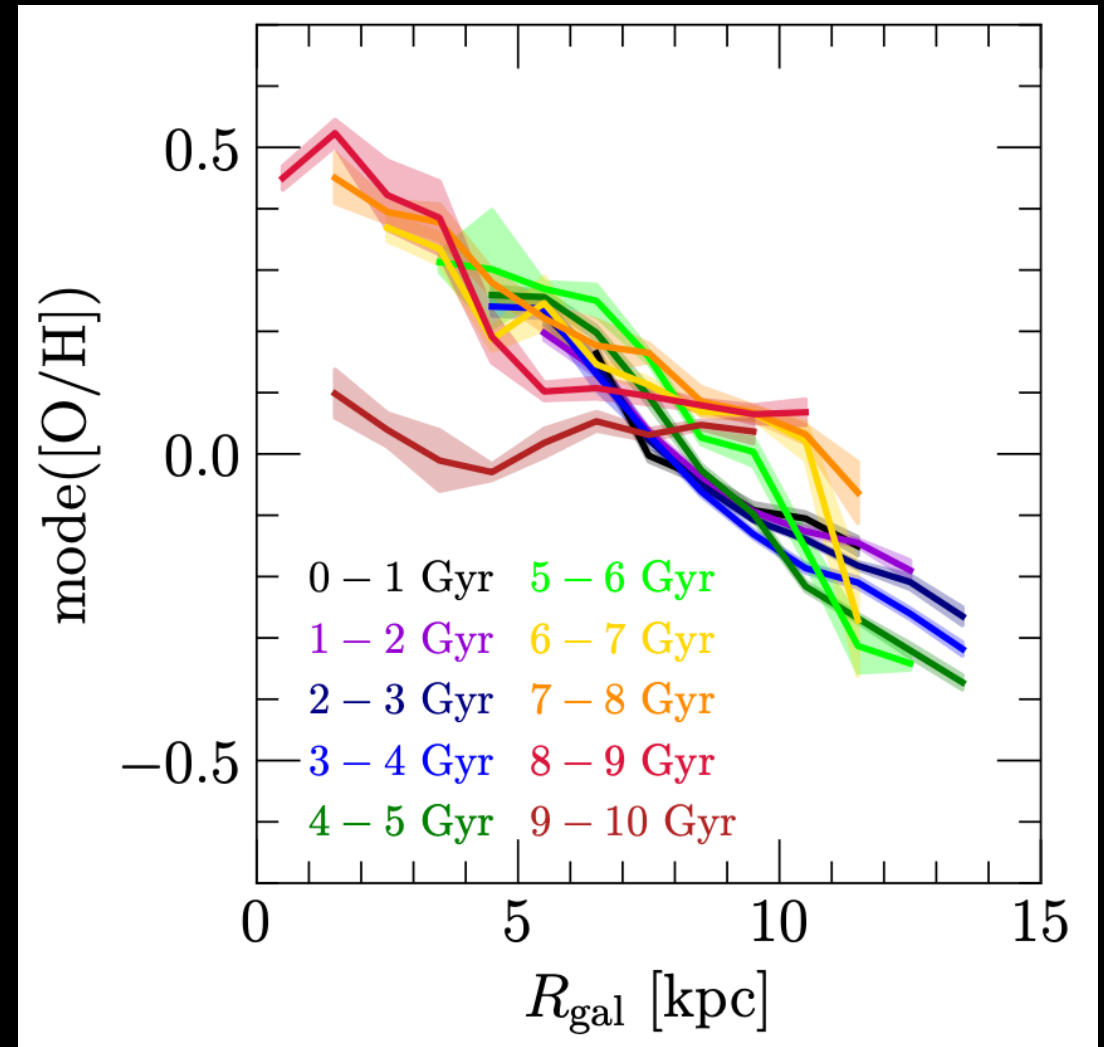
The ISM evolves in many ways, but metal-hydrogen mixture is unaffected



An Apparent Equilibrium in the Milky Way

Shockingly little variation in metallicity with population age

- Appears to exist across the Galaxy
- Reached ~ 9 Gyr ago in this sample (APOGEE red giants)

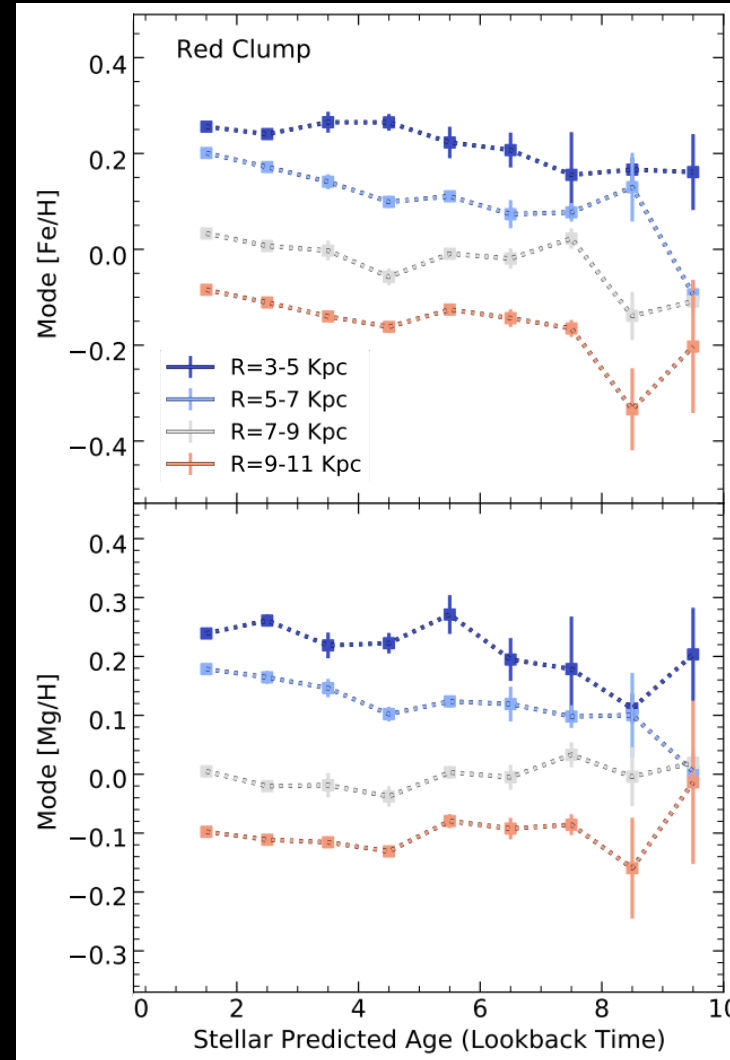


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Jack Roberts
(Ohio State)

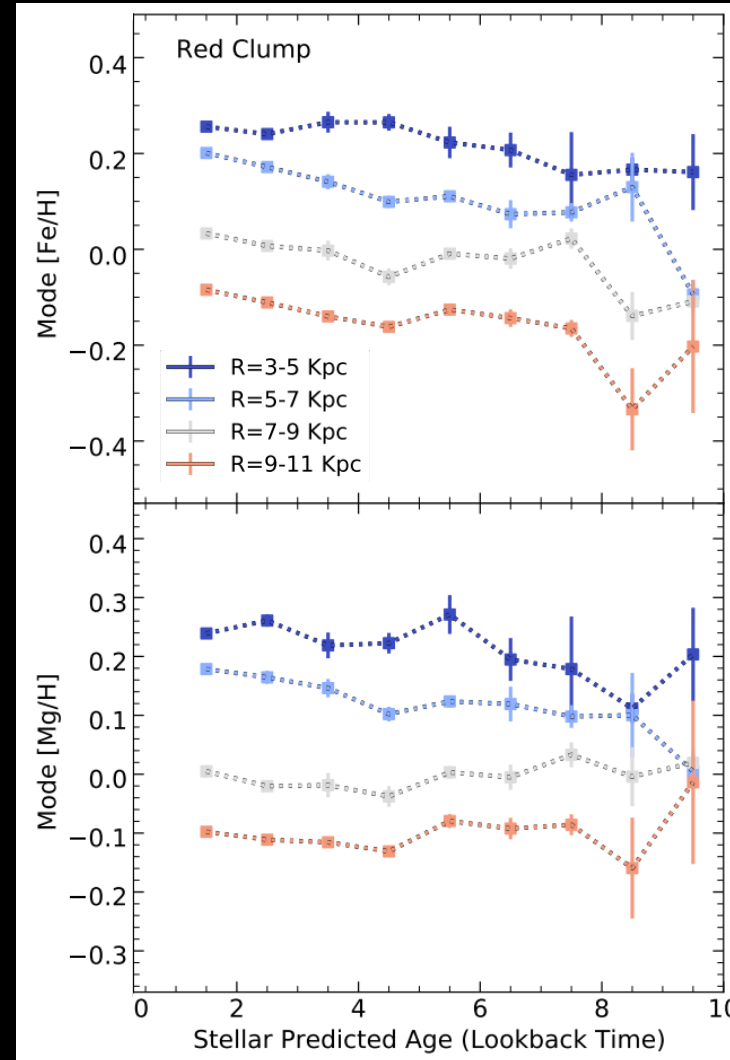
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- Magrini et al. (2022)
- Willett et al. (2023)
- Gallart et al. (2024)



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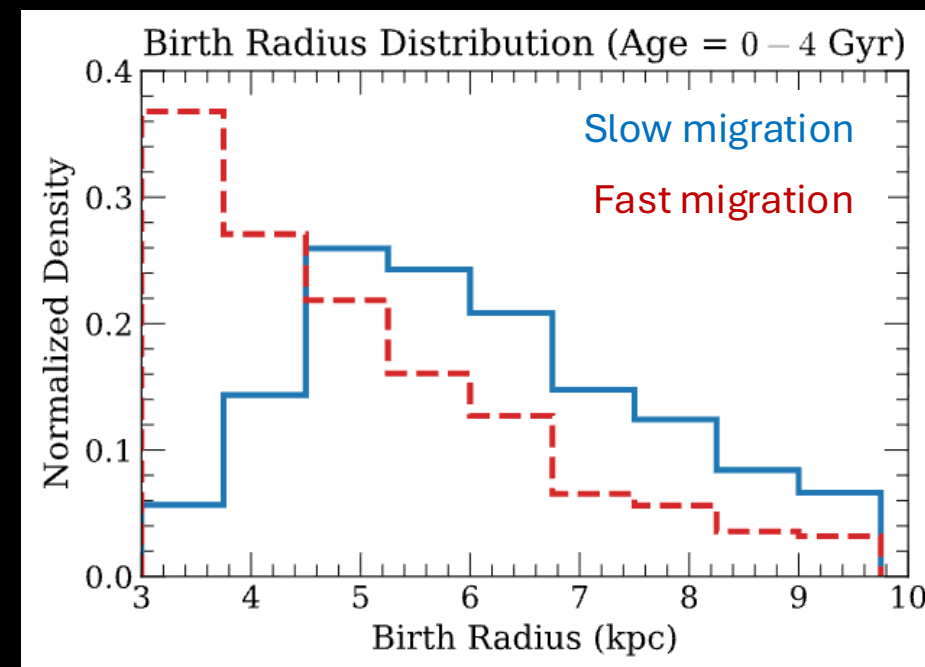
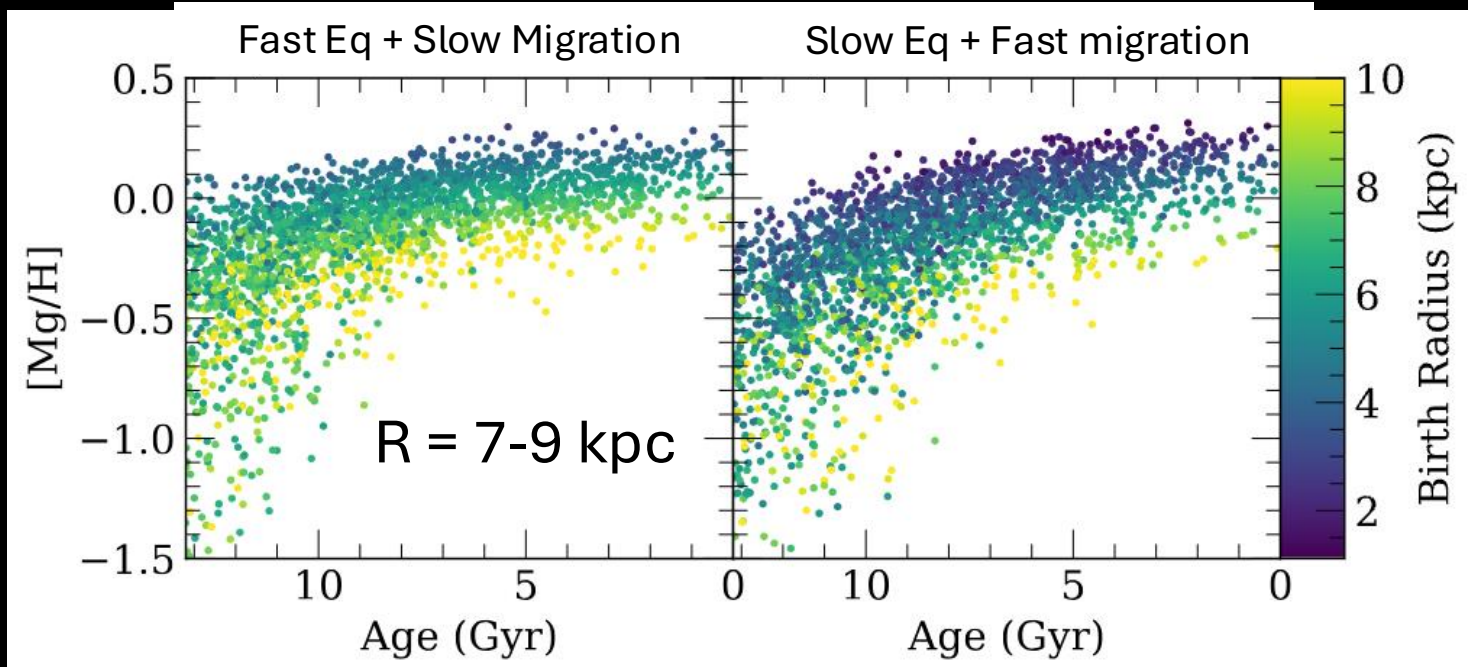
Alternatively: Fast Radial Migration

Local stars no longer reflect local enrichment history

- Roškar et al. (2008)
- Loebman et al. (2011)



Christopher Giudice
(San Francisco State)



Chemical Equilibrium from Radial Gas Flows?

New coefficient describing rate relative to SFR:

$$Z \rightarrow Z_{eq} = \frac{y}{1 + \eta - r}$$

Chemical Equilibrium from Radial Gas Flows?

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$$\mu_Z \equiv \dot{\Sigma}_{Z,flow} / \dot{\Sigma}_\star$$

Chemical Equilibrium from Radial Gas Flows?

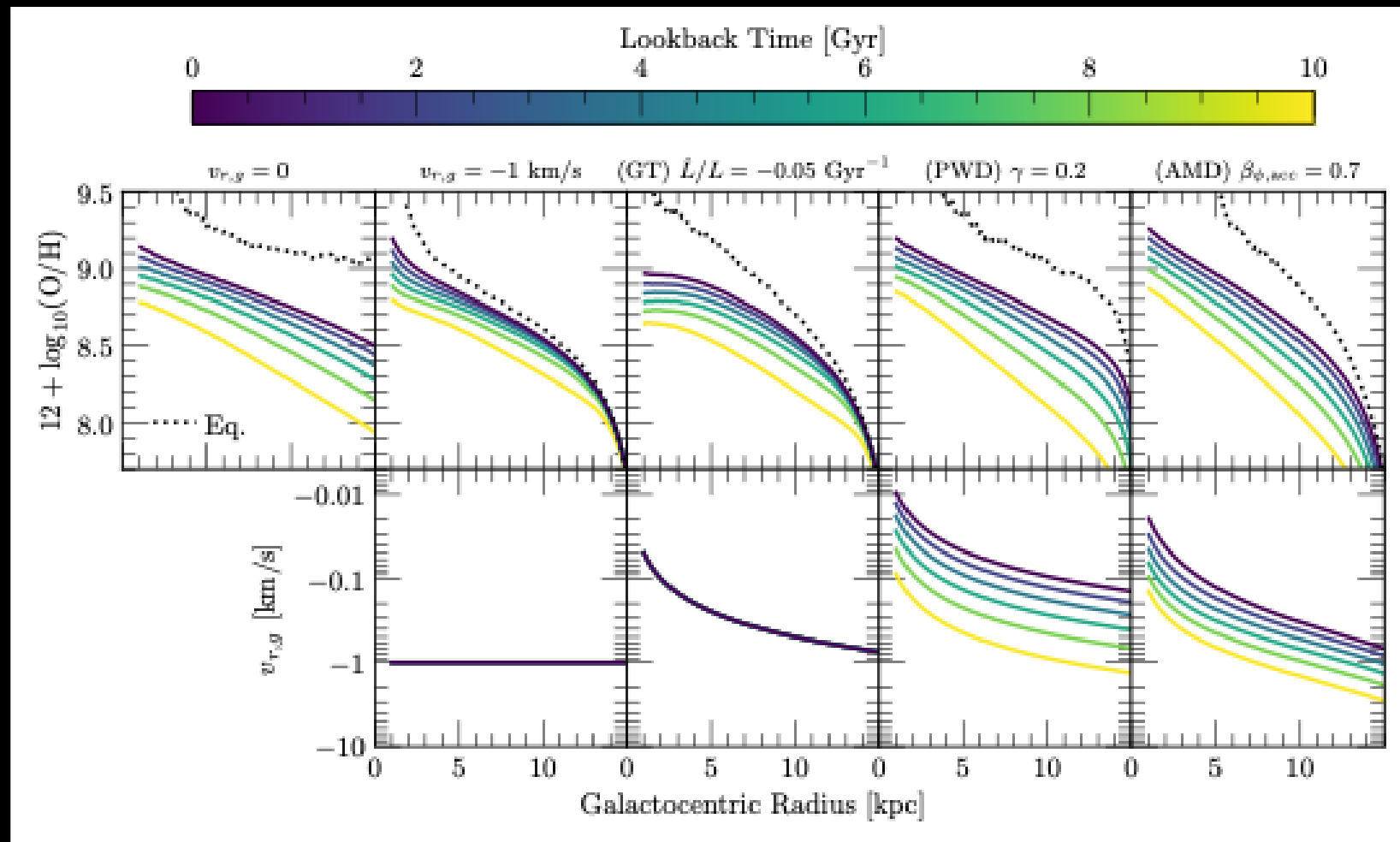
New coefficient describing rate relative to SFR:

$$Z \rightarrow Z_{eq} = \frac{y}{1 + \eta - \mu_Z - r}$$

$$\eta > 0 \Rightarrow \dot{M}_{out} > 0 \text{ (removing mass)}$$

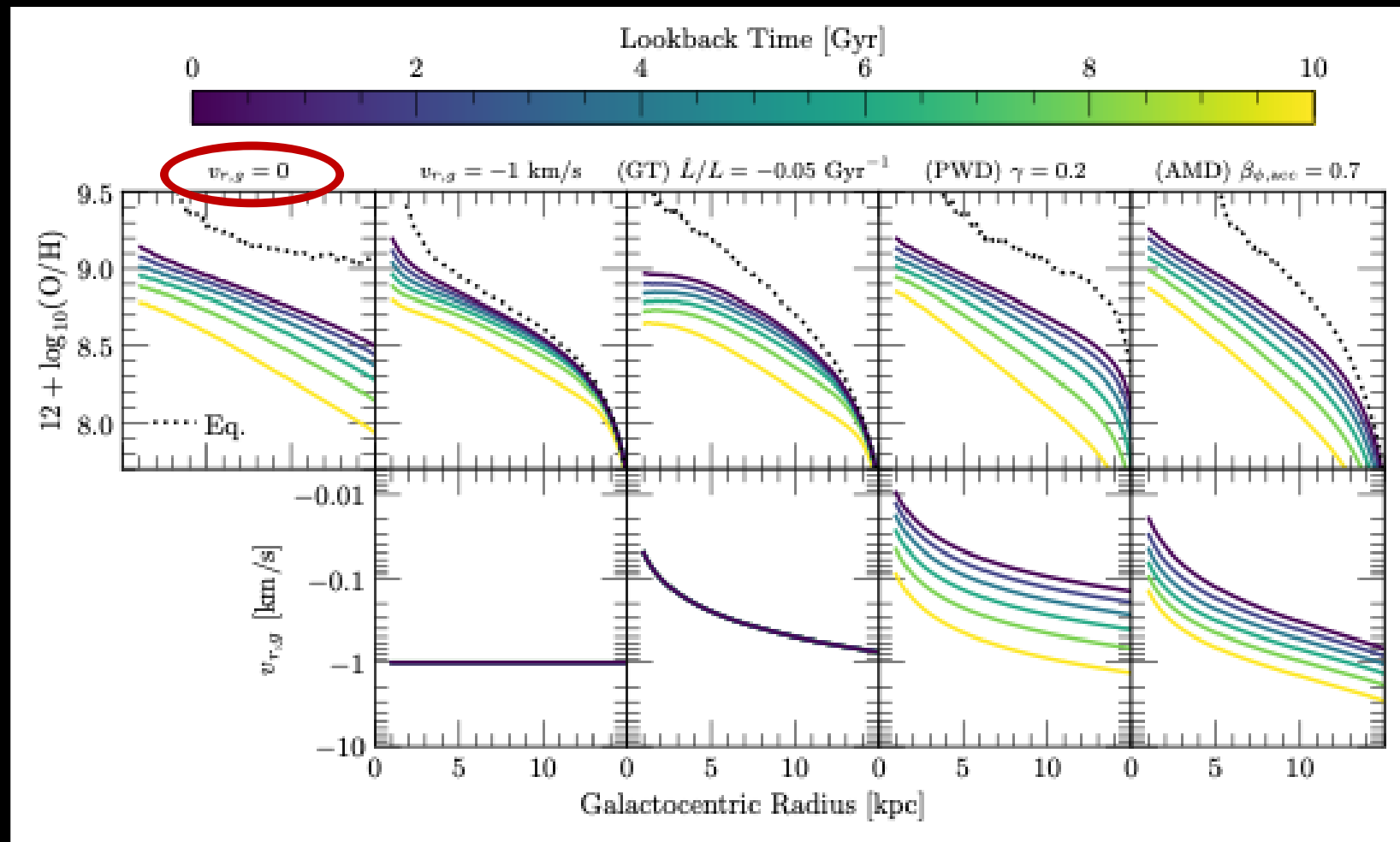
$$\mu > 0 \text{ when } \dot{\Sigma}_{flow} > 0 \text{ (adding mass)}$$

Four Radial Flow Prescriptions



Four Radial Flow Prescriptions

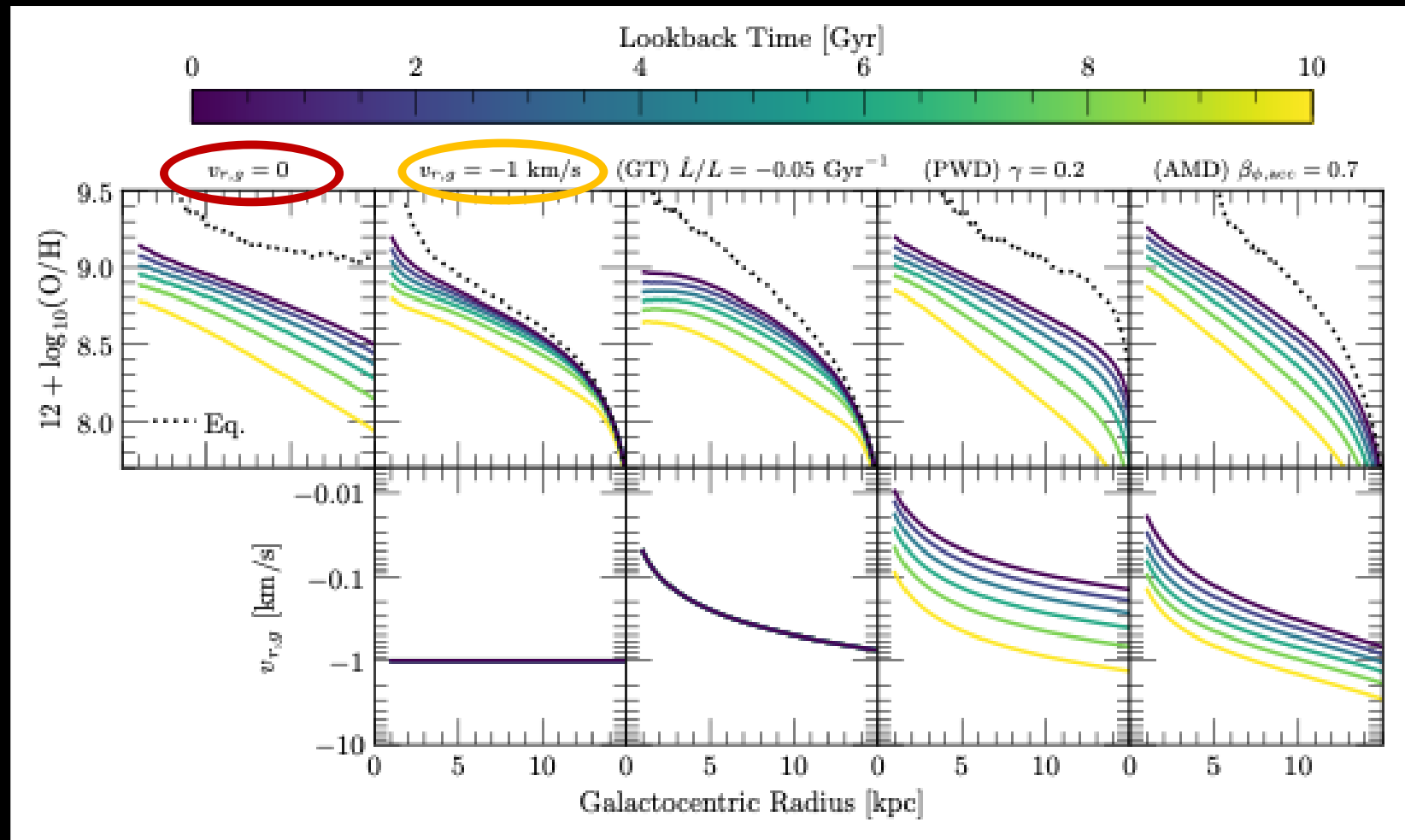
No Flow



Four Radial Flow Prescriptions

No Flow

Constant Velocity

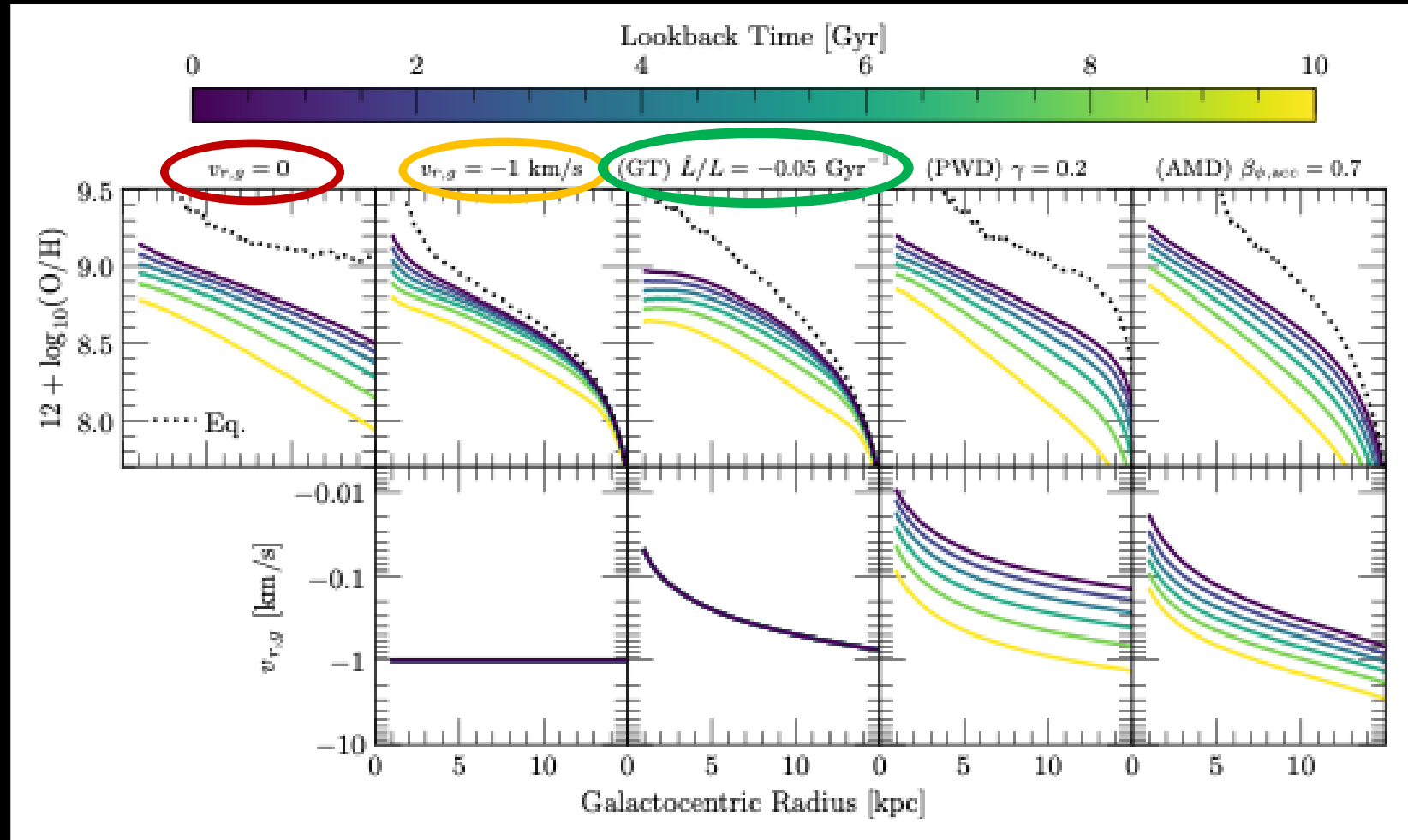


Four Radial Flow Prescriptions

No Flow

Constant Velocity

Global Torque: constant
angular momentum loss



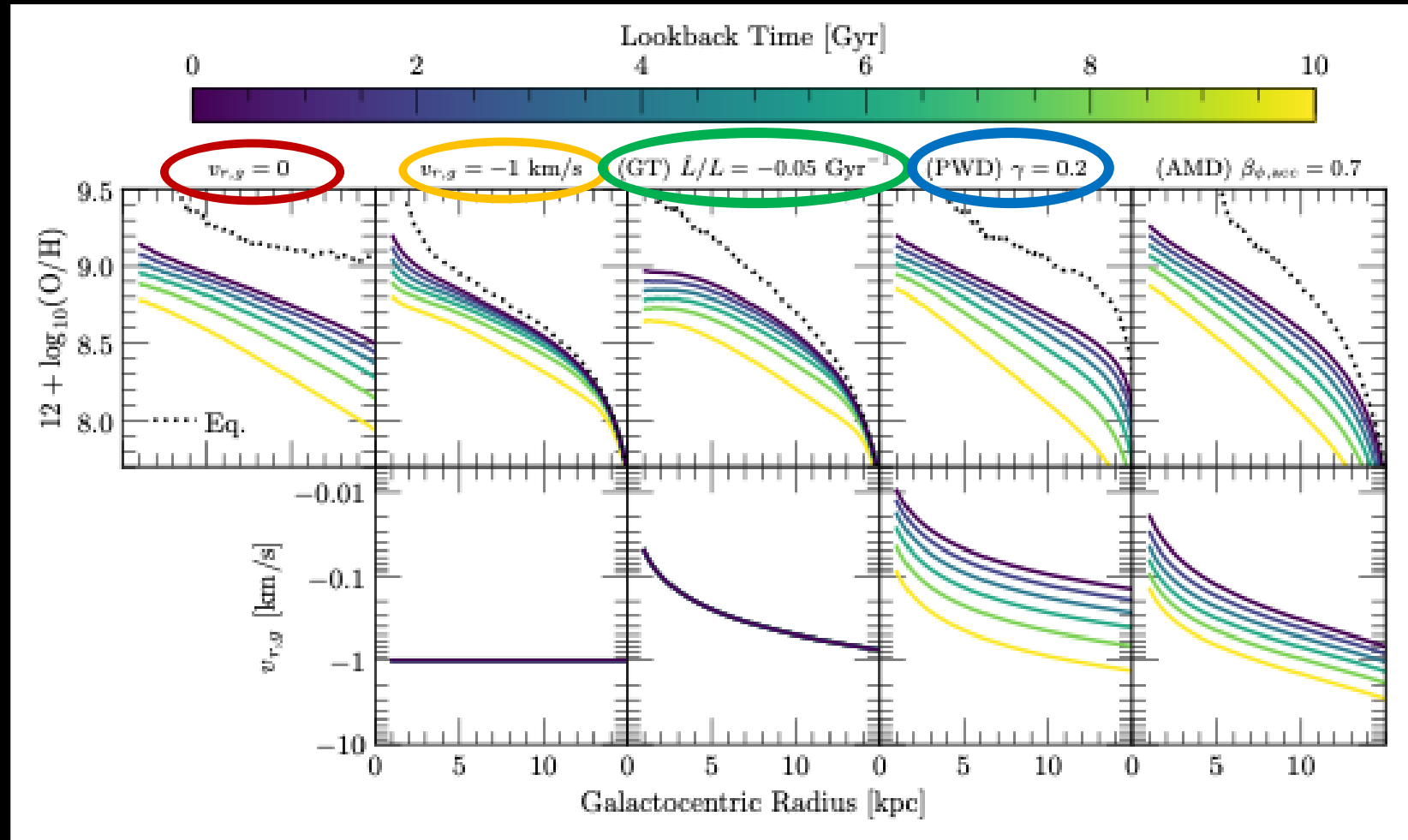
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Circular velocity grows



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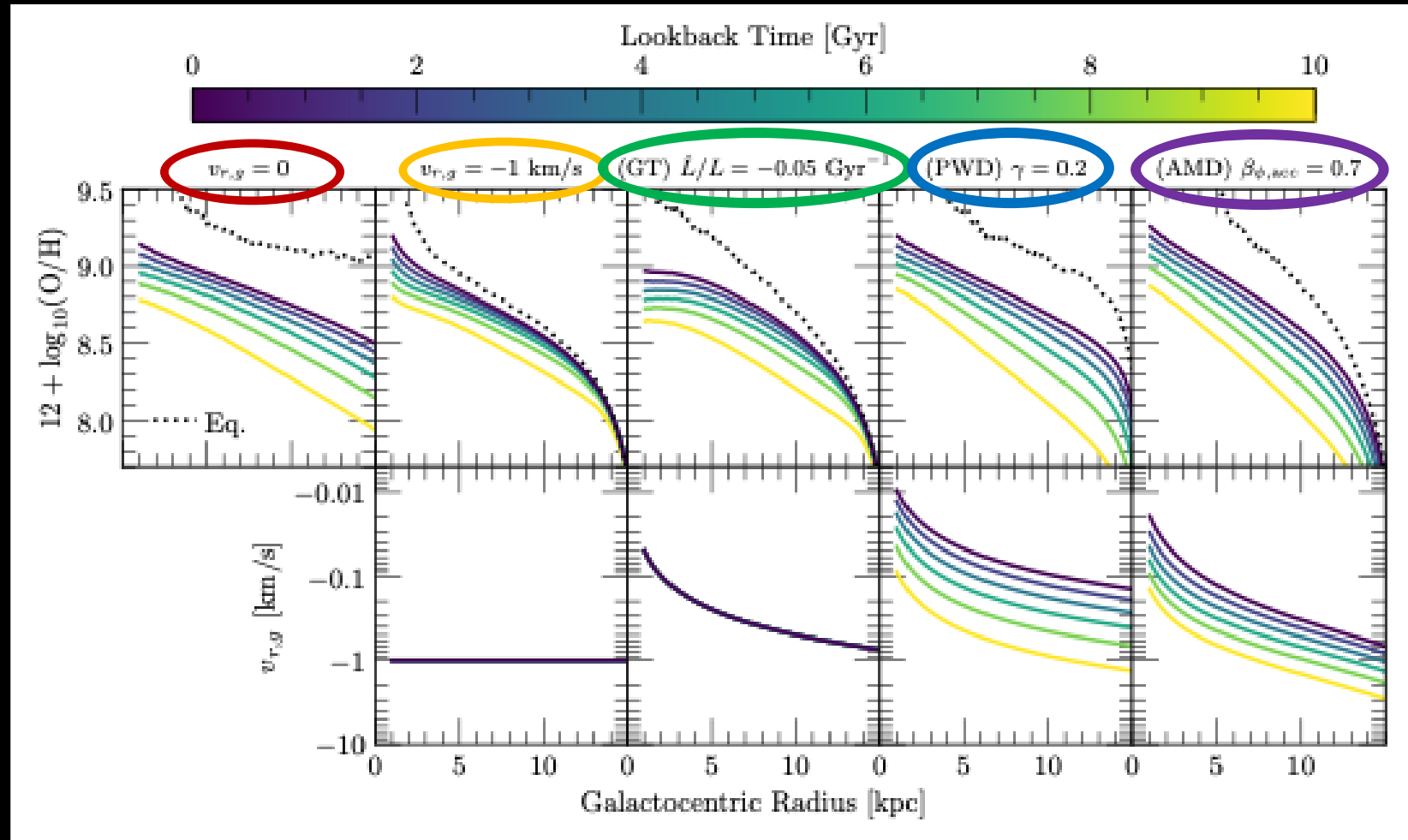
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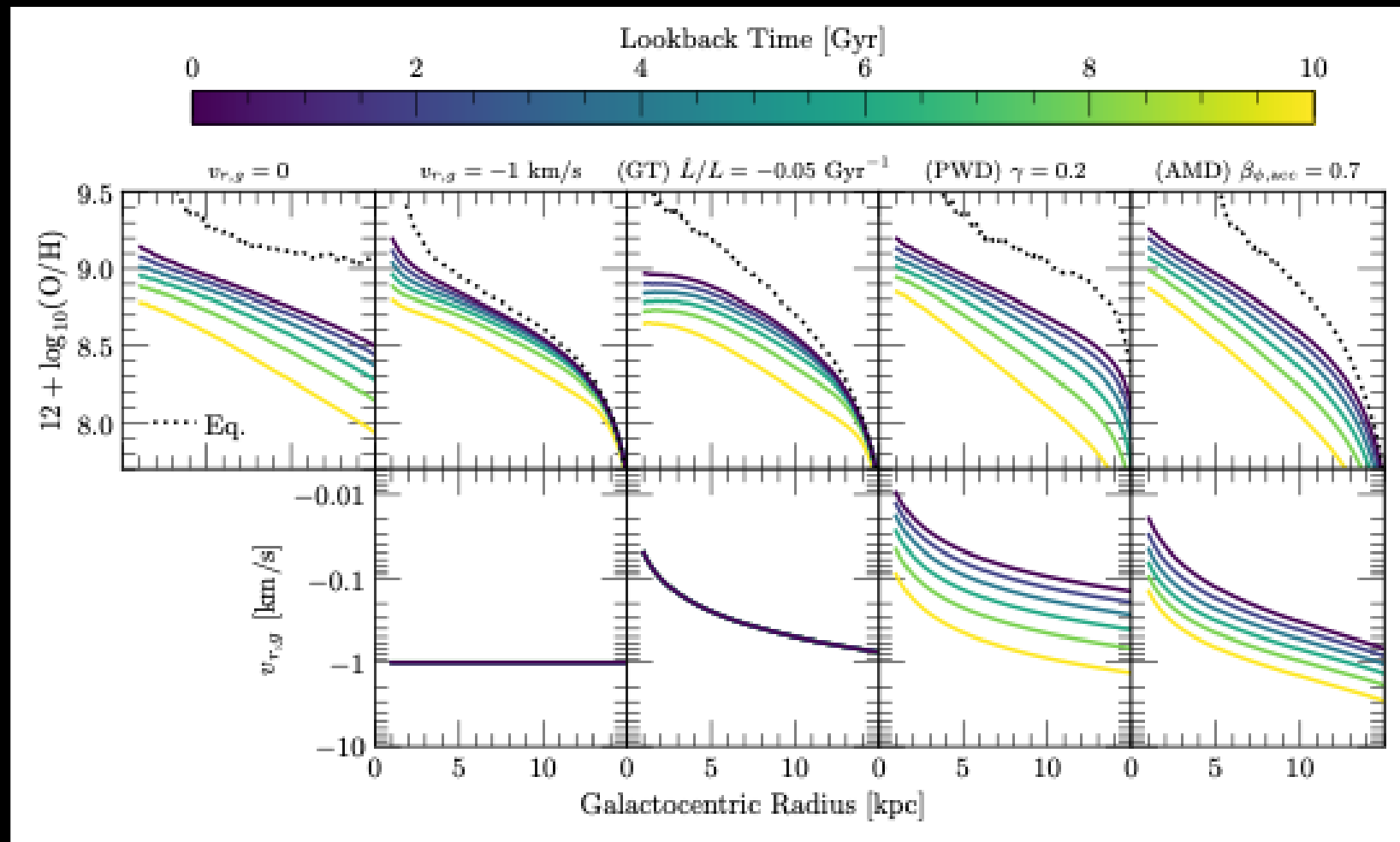
Global Torque: constant
angular momentum loss

Potential Well Deepening:
Circular velocity grows

Angular Momentum Dilution:
 $L_{z,acc} < L_{z,ISM}$

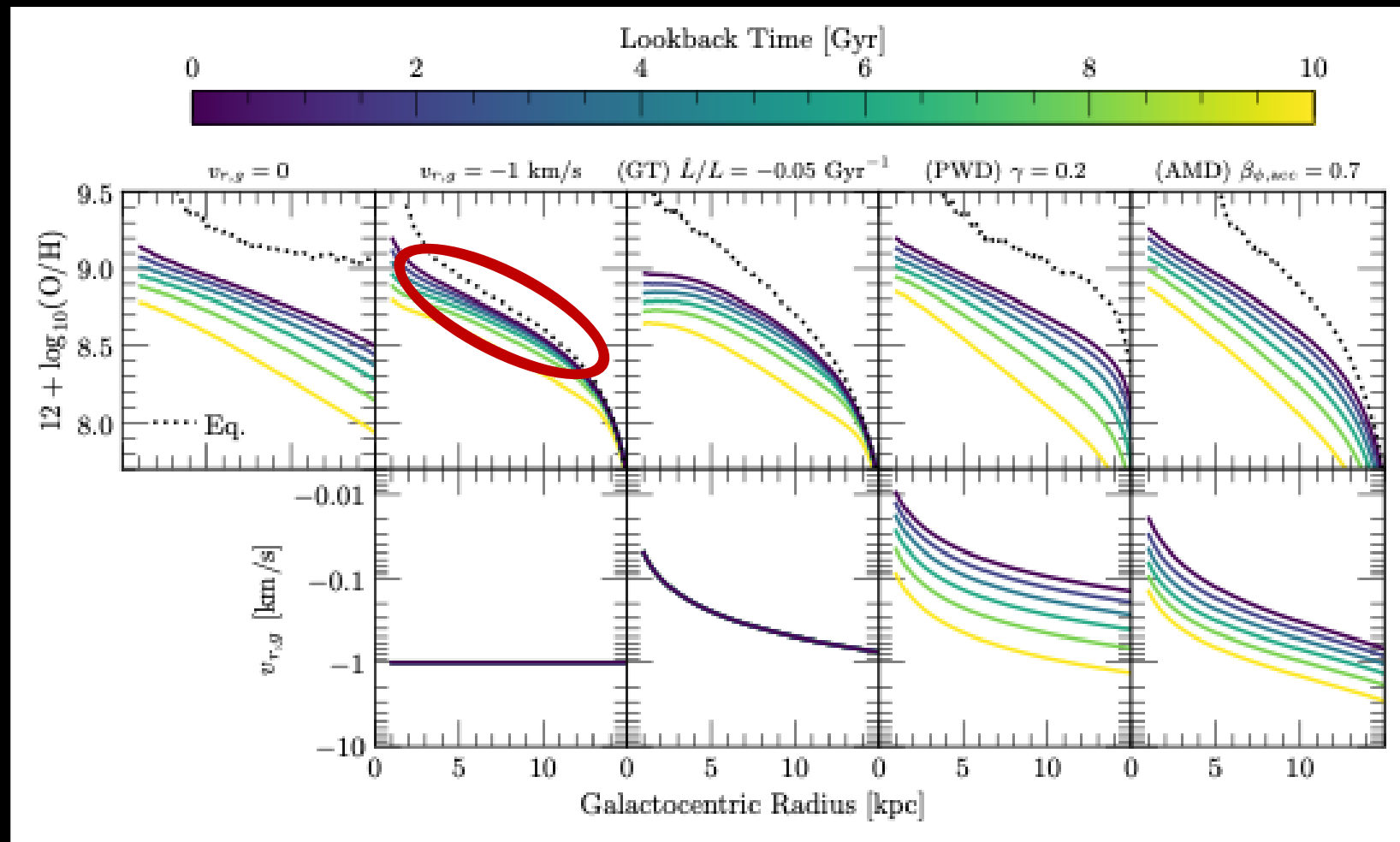


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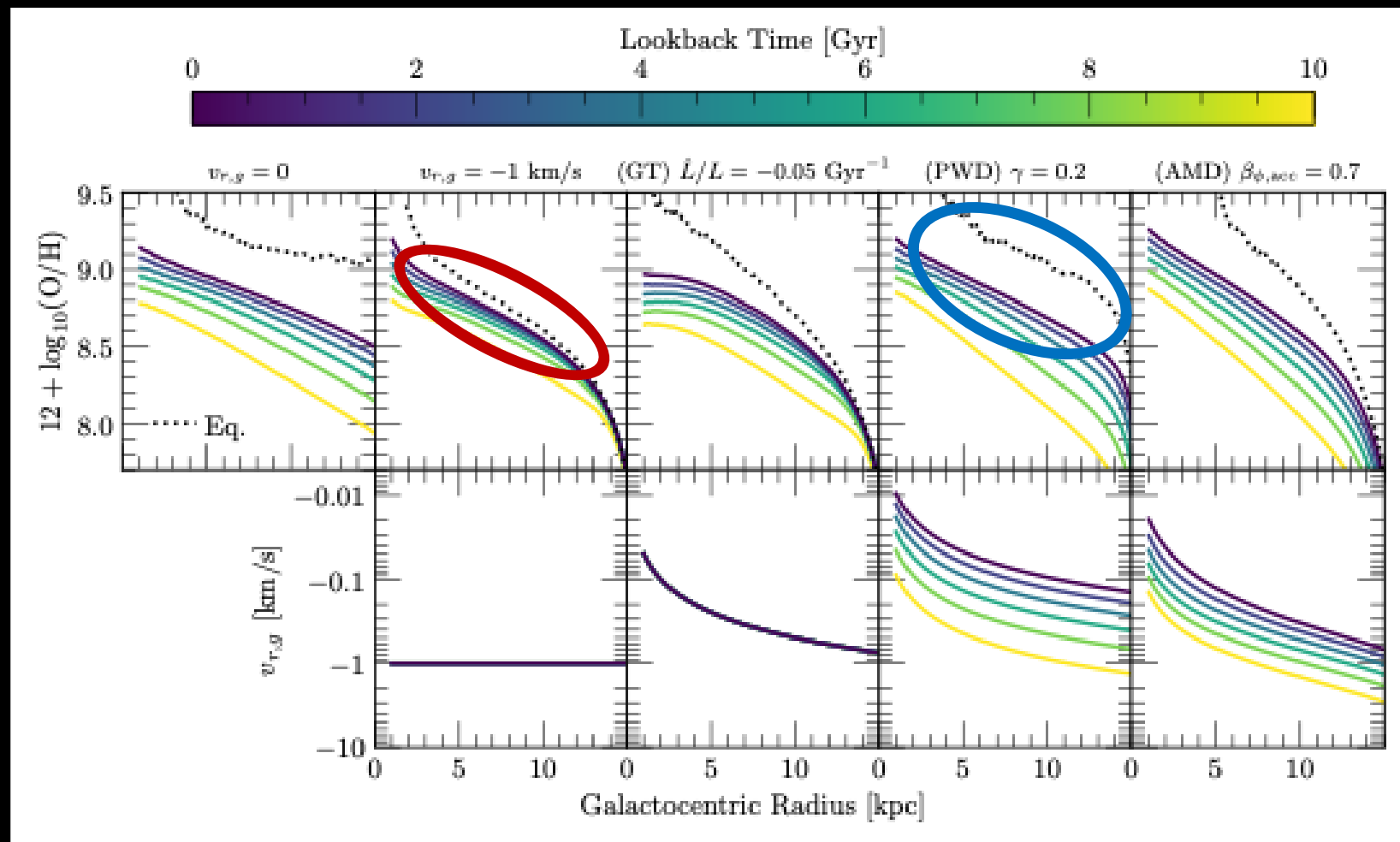
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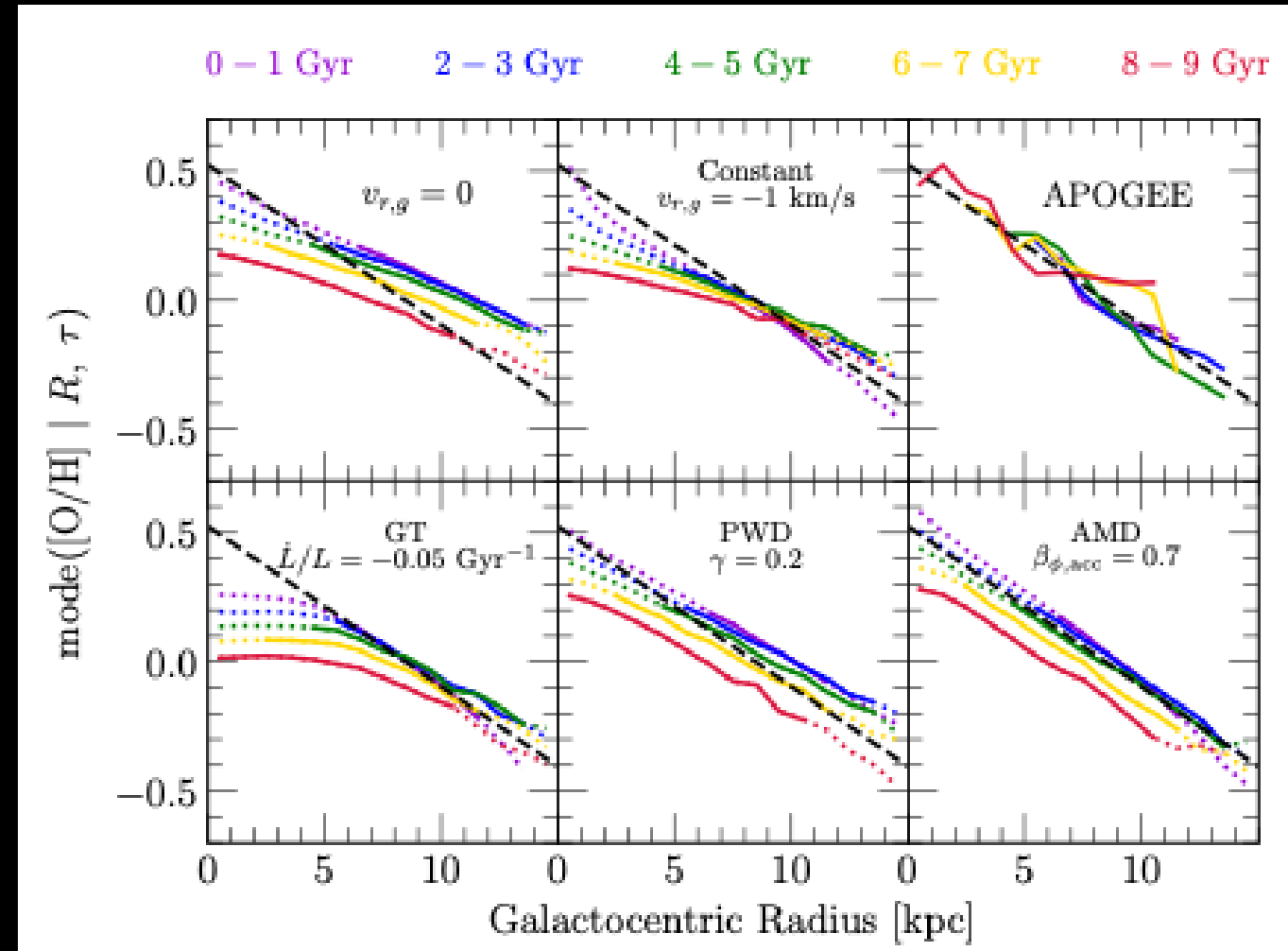
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Age-Metallicity Relation with Radial Gas Flows

All models improve agreement with data by steepening gradient

Cannot wipe out age trend entirely

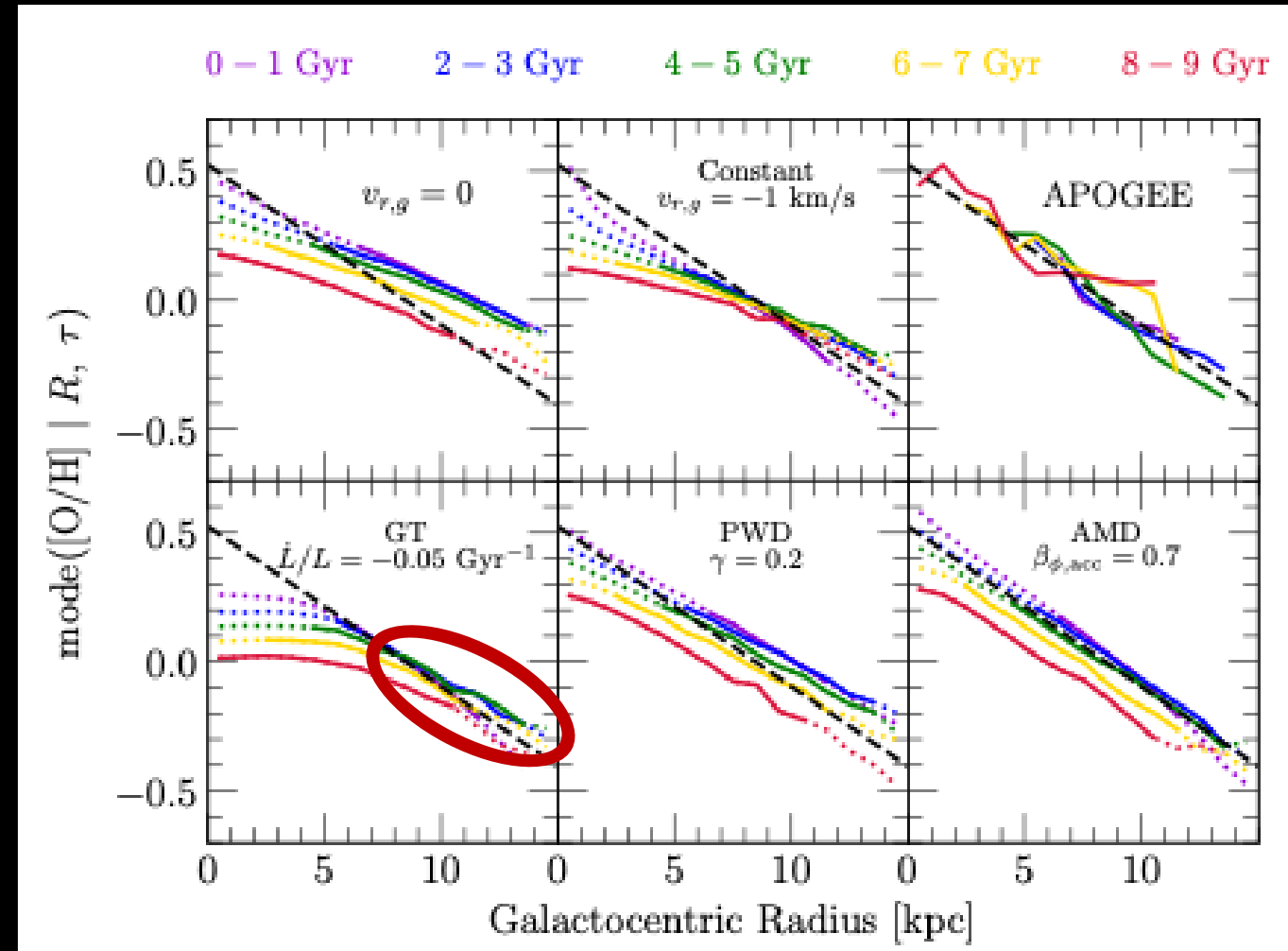


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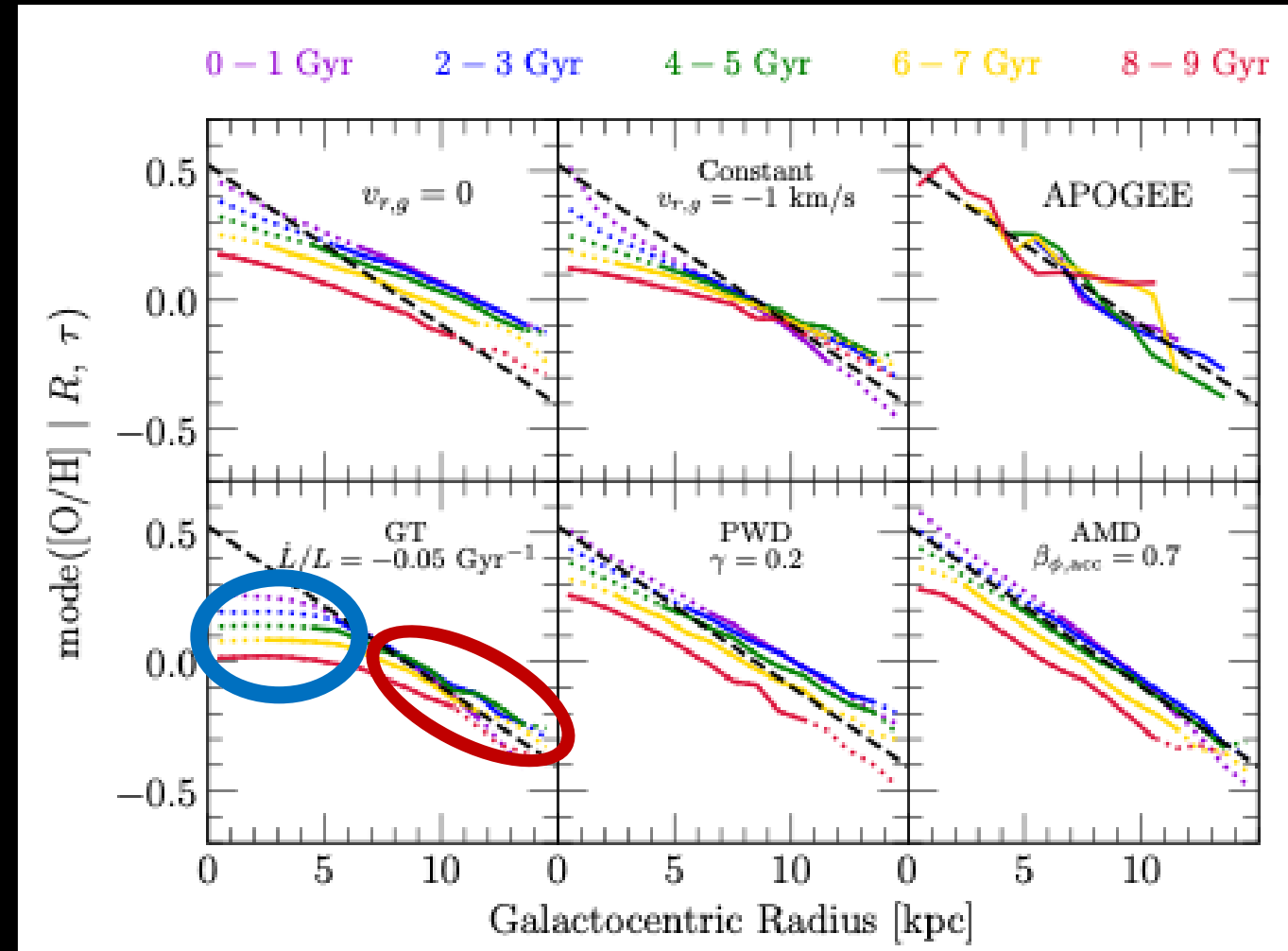


Age-Metallicity Relation with Radial Gas Flows

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- Equilibrium region
- Disequilibrium region



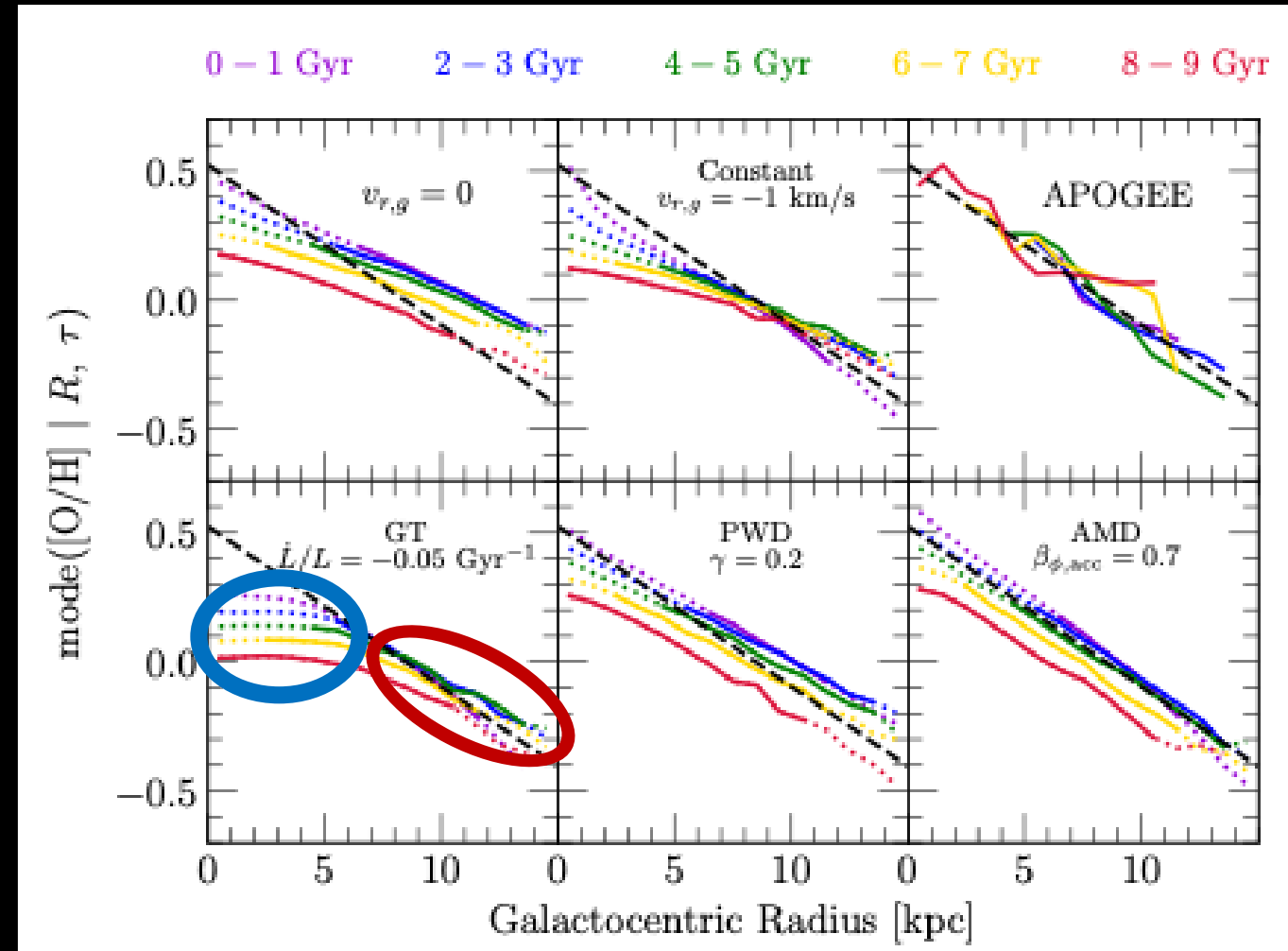
Age-Metallicity Relation with Radial Gas Flows

Mass Conservation:

$$\int_0^\infty \mu Z \dot{\Sigma}_* R dR = 0$$

$\mu < 0 \Rightarrow$ Equilibrium

$\mu > 0 \Rightarrow$ Disequilibrium



Hydrogen sets the Metallicity Scale

Outflows and radial flows $\rightarrow$ adjust amount of H that mixes with new metals

$$Z \rightarrow Z_{eq} = \frac{y}{1 + \eta - \mu_Z - r - \tau_{\star}/\tau_{sfh}}$$

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How many stars form?

... Per accreted baryon?

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$\dot{Z}$

$\dot{H}$

Hydrogen sets the Metallicity Scale

Early lesson learned from GCE: galaxies are not closed boxes

Outflows & mass-metallicity relation: dwarf galaxies need more H to form the same amount of stars

If equilibrium scenario is accurate: H does not hang around in a galaxy for long before forming stars, being ejected, moving in a flow, etc.

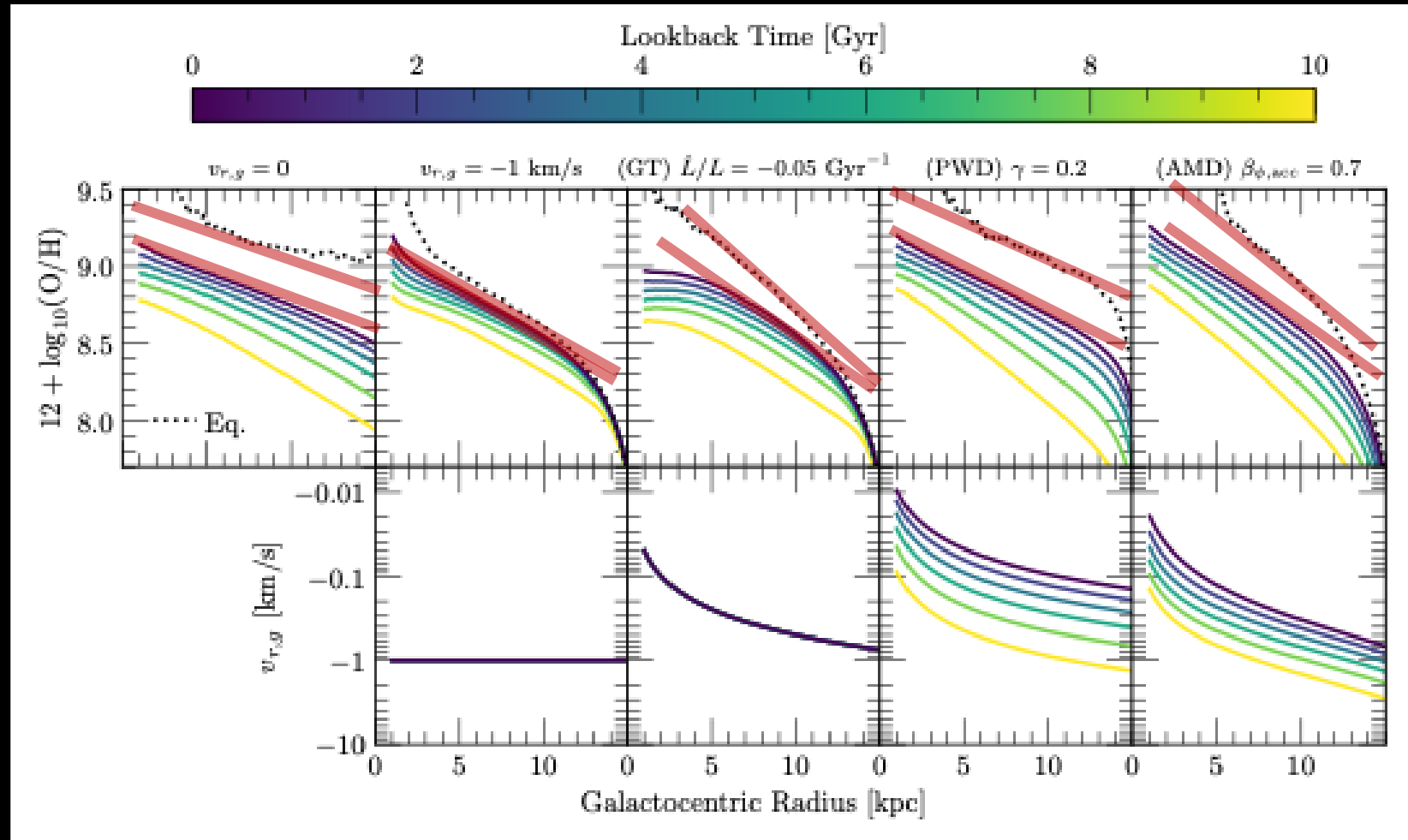
Metallicity gradients: same quantities – now R-dependence within galaxies

Gradient Evolution as an Equilibration Process

$$\frac{\partial \ln Z_{ISM}}{\partial R} \approx \frac{\partial \ln Z_{eq}}{\partial R}$$

even when

$$Z_{ISM} \ll Z_{eq}$$



Allows Analytic Approximations of $\nabla[Z/H]$

$$\frac{\partial}{\partial R} \left[1 + \eta - \mu - r - \frac{\tau_{\star}}{\tau_{sfh}} \right] \rightarrow \dots \rightarrow \dots \rightarrow \nabla \ln Z_{eq} \sim \nabla \ln Z_{ISM}$$

Folding in disequilibrium corrections, accurate to ~ 0.01 dex/kpc

